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Article

# The Photon Energy Density Parameter $\Omega_\gamma$ of the Universe Exactly Derived

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**Abstract:** We will demonstrate that the photon energy density parameter  $\Omega_\gamma$  of the universe can be derived exactly in  $R_H = ct$  cosmology. We find that it must be  $\Omega_\gamma = \frac{1}{5760\pi} \approx 5.52621330180192 \times 10^{-5}$ . This indicates that there is no uncertainty in the radiation density within at least some sub-classes of  $R_H = ct$  cosmology.

**Keywords:** photon energy density parameter; Friedmann equation; critical density; thermodynamical Friedmann equation

## Deriving the Photon Energy Density Parameter $\Omega_\gamma$

The photon energy density:  $\rho_\gamma$  is normally found by doing the following integral:

$$\rho_{\gamma,0}c^2 = \int_0^\infty h\nu n(\nu)d\nu = a_b T_0^4 \quad (1)$$

which mean we have:

$$\rho_{\gamma,0} = \frac{a_b T_0^4}{c^2} \quad (2)$$

where the radiation density constant:  $a_b$  is given by:

$$a_b = \frac{8\pi^5 k_b^4}{15c^3 h^3} = \frac{4}{c} \sigma \quad (3)$$

In 1978, Emslie and Green [1] presented the photon energy density relative to the critical Friedmann energy density as:

$$\begin{aligned} \Omega_{\gamma,0} &= \frac{\rho_\gamma}{\rho_c} \\ \Omega_{\gamma,0} &= \frac{\frac{a T_0^4}{c^2}}{\frac{3H_0^2}{8\pi G}} \end{aligned} \quad (4)$$

Haug and Tatum [2] have recently shown that the Friedmann [3] equation can be expressed in thermodynamic form, indicating that the critical density must be given by:

$$\rho_{c,0} = \frac{3H_0}{8\pi G} = T_0^4 \frac{23040\pi}{c^3} \sigma \quad (5)$$

By replacing this into equation (4) one get:

$$\begin{aligned} \Omega_{\gamma,0} &= \frac{\frac{a T_0^4}{c^2}}{T_0^4 \frac{23040\pi}{c^3} \sigma} \\ \Omega_{\gamma,0} &= \frac{4\sigma}{23040\pi\sigma} \\ \Omega_{\gamma,0} &= \frac{1}{5760\pi} \approx 5.52621330180192 \times 10^{-5} \end{aligned} \quad (6)$$

This implies that the photon density parameter remains exact for all cosmic time epochs. This consistency aligns with at least two potential types of  $R_h = ct$  cosmological models. It is conceivable within growing black hole models, as exemplified in [4], and also within steady-state black hole cosmology, where the metric from general relativity suggests that density inside the black hole varies as one approaches the center, as discussed in [5,6]. Additionally, this holds true in the extremal universe scenario, where the density inside the black hole's Hubble sphere varies, as indicated in [7]. The idea of black hole cosmology goes at least back to 1972 by Pathria [8] and is actively discussed to this day [9–17], even if the  $\Lambda$ -CDM model currently takes most headlines.

When observing the black hole (Hubble sphere), it resembles a growing black hole model due to the density change along the radius in the steady-state black hole.

In  $R_h = ct$  black hole cosmology in earlier epochs of the universe, we must have:

$$\rho_{\gamma,t} = \frac{a_b T_0^4 (1+z)^4}{c^2} = \frac{a_b T_t^4}{c^2} \quad (7)$$

Furthermore, we must have a critical density, as demonstrated by Haug and Tatum [2,18], of:

$$\rho_{c,t} = \frac{3H_t}{8\pi G} = T_0^4 (1+z)^4 \frac{23040\pi}{c^3} \sigma = T_t^4 \frac{23040\pi}{c^3} \sigma \quad (8)$$

By substituting these into the equation below, we obtain:

$$\begin{aligned} \Omega_{\gamma,t} &= \frac{\rho_{\gamma}}{\rho_{c,t}} \\ \Omega_{\gamma,t} &= \frac{\frac{a T_t^4}{c^2}}{T_t^4 \frac{23040\pi}{c^3} \sigma} \\ \Omega_{\gamma,t} &= \frac{4\sigma}{23040\pi\sigma} \\ \Omega_{\gamma,t} &= \frac{1}{5760\pi} \approx 5.52621330180192 \times 10^{-5} \end{aligned} \quad (9)$$

In a black hole growing  $R_h = ct$  cosmology during the entire cosmic epoch, the photon radiation density ratio remains constant and exact. This is likely inconsistent with the predictions of the  $\Lambda$ -CDM model in earlier times, but it should be correct for the current time as well.

**Conflicts of Interest:** The author declare no conflict of interest.

**Data Availability Statement:** No data was used for this study.

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