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Posted Date: 29 November 2023

doi: 10.20944/preprints202311.1902.v1

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Article

Symmetry and Asymmetry of One-Way Einstein-Rosen Bridge inside Black Hole

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Abstract: We elucidate the versatility of intriguing symmetry and asymmetry phenomena outside and inside black holes. For description of the test particle (planet and photon) motion it is used the Kerr-Newman metric of the rotating and electrically charged black hole. In particular, it is demonstrated the symmetry and asymmetry of the one-way Einstein-Rosen bridge inside black hole toward and through the plethora of endless and infinite universes.

Keywords: General Relativity; black holes; black hole shadow

PACS: 04.70.Bw; 98.35.Jk; 98.62.Js

1. Introduction

We elucidate the versatility of intriguing symmetry and asymmetry phenomena outside and inside black holes. For description of the test particle (planet and photon) motion it is used the Kerr-Newman metric of the rotating and electrically charged black hole. In particular, it is demonstrated the symmetry and asymmetry of the one-way Einstein-Rosen bridge inside black hole toward and through the plethora of endless and infinite universes.

2. Kerr-Newman Metric for Rotating and Electrically Charged Black Hole

The famous Kerr-Newman metric or geometry (see e. g., [1–5]), which is the exact solution of Einstein's equations [6–10] for a rotating and electrically black hole, is:

$$ds^2 = -\frac{\Delta}{\Sigma} [dt - a \sin^2 \theta d\varphi]^2 + \frac{\sin^2 \theta}{\Sigma} [(r^2 + a^2)d\varphi - a dt]^2 + \frac{\Sigma}{\Delta} dr^2 + \Sigma d\theta^2, \quad (1)$$

where (r, θ, ϕ) are spherical coordinates and t is the time of static distant observer at the asymptotically radial infinity. In this metric

$$\Delta = r^2 - 2Mr + a^2 + q^2, \quad \Sigma = r^2 + a^2 \cos^2 \theta, \quad (2)$$

M — black hole mass, q — black hole electric charge, $a = J/M$ — specific black hole angular momentum (spin). The two roots of equation $\Delta = 0$ are $r_+ = 1 + \sqrt{M^2 - a^2 - q^2}$ — the event horizon of the black hole and $r_- = 1 - \sqrt{M^2 - a^2 - q^2}$ — the Cauchy horizon of the black hole.

For simplification of equation and presentation of Figures we will use often units $G = 1$, $c = 1$, $M = 1$ and dimensionless radius $r \Rightarrow r/M$ and dimensionless time $t \Rightarrow t/M$.

In the Kerr-Newman metric there the following integrals of motion for test particles [2]: μ — particle mass, E — particle total energy, L — particle azimuthal angular moment and Q — the Carter

constant, related with the non-equatorial particle motion. The corresponding equations of test particle motion in the Kerr-Newman metric in the differential form are:

$$\Sigma \frac{dr}{d\tau} = \sqrt{R}, \quad (3)$$

$$\Sigma \frac{d\theta}{d\tau} = \sqrt{\Theta}, \quad (4)$$

$$\Sigma \frac{d\varphi}{d\tau} = -a(aE - \frac{L}{\sin^2 \theta}) + \sin^2 \theta \frac{a}{\Delta}, \quad (5)$$

$$\Sigma \frac{dt}{d\tau} = -a(aE \sin^2 \theta - L) + (r^2 + a^2) \frac{P}{\Delta}. \quad (6)$$

Here $P = E(r^2 + a^2) - aL$, τ — the proper time of a test massive particle or an affine parameter along the trajectory of a massless particle ($\mu = 0$) like photon. Respectively, the effective radial potential $R(r)$ is

$$R(r) = P^2 - \Delta[\mu^2 r^2 + (L - aE)^2 + Q], \quad (7)$$

and the effective polar potential $\Theta(\theta)$ is

$$\Theta(\theta) = Q - \cos^2 \theta [a^2(\mu^2 - E^2) + L^2 \sin^{-2} \theta]. \quad (8)$$

Trajectories of massive particles ($\mu \neq 0$) depend on three parameters: $\gamma = E/\mu$, $\lambda = L/\mu$ and Q/μ^2 . Meantime, trajectories of massless particles like photons (null geodesics) depend on two parameters: λ and Q .

The nontrivial specific feature of the rotating Kerr black hole ($a \neq 0$) is the existence of so-called *ergosphere* [4,5,9,11,12] with the outer boundary

$$r_{\text{ES}}(\theta) = 1 + \sqrt{1 - q^2 - a^2 \cos^2 \theta}. \quad (9)$$

Inside the ergosphere any test object is dragged into insuperable rotation around black hole with infinite azimuthal winding by approaching the black hole event horizon.

In the following Sections we will describe the symmetry and asymmetry of test object motion in the gravitational field of the Kerr-Newman black hole. We use equations of motion in the Kerr-Newman metric (3)–(6) in our analytic and numerical calculations of test particle geodesic trajectories [13–23].

3. One-Way Einstein-Rosen Bridge Inside Black Hole

We start to elucidate the versatility of intriguing symmetry and asymmetry phenomena outside and inside black holes by using the Carter-Penrose diagrams (for details see, i. e., [4,5,9,10]), describing in particular, the global space-time structure of black holes. The evident manifestation of *symmetry* of this global structure is infinite space volumes as outside and inside the black hole event horizon. See at Figure 1 the corresponding Carter-Penrose diagram for the Reissner-Nordström black hole, which is a special spherically *symmetric* case of Kerr-Newman black hole without rotation i. e., $a = 0$ but $q \neq 0$. From the pure geometric point of view this diagram is both left-right and up-down symmetric. On the contrary, from the physical or space-time point of view this diagram is absolutely *asymmetric* due to the inexorable upward flow of time not only at this diagram but throughout the whole universe. More precisely it means that in the General Relativity all objects are allowed to move only inside the the upward directed light cones (at $\pm 45^\circ$ with respect to the upward direction. The upward directed light cone is the inexorable *asymmetry* of the world.

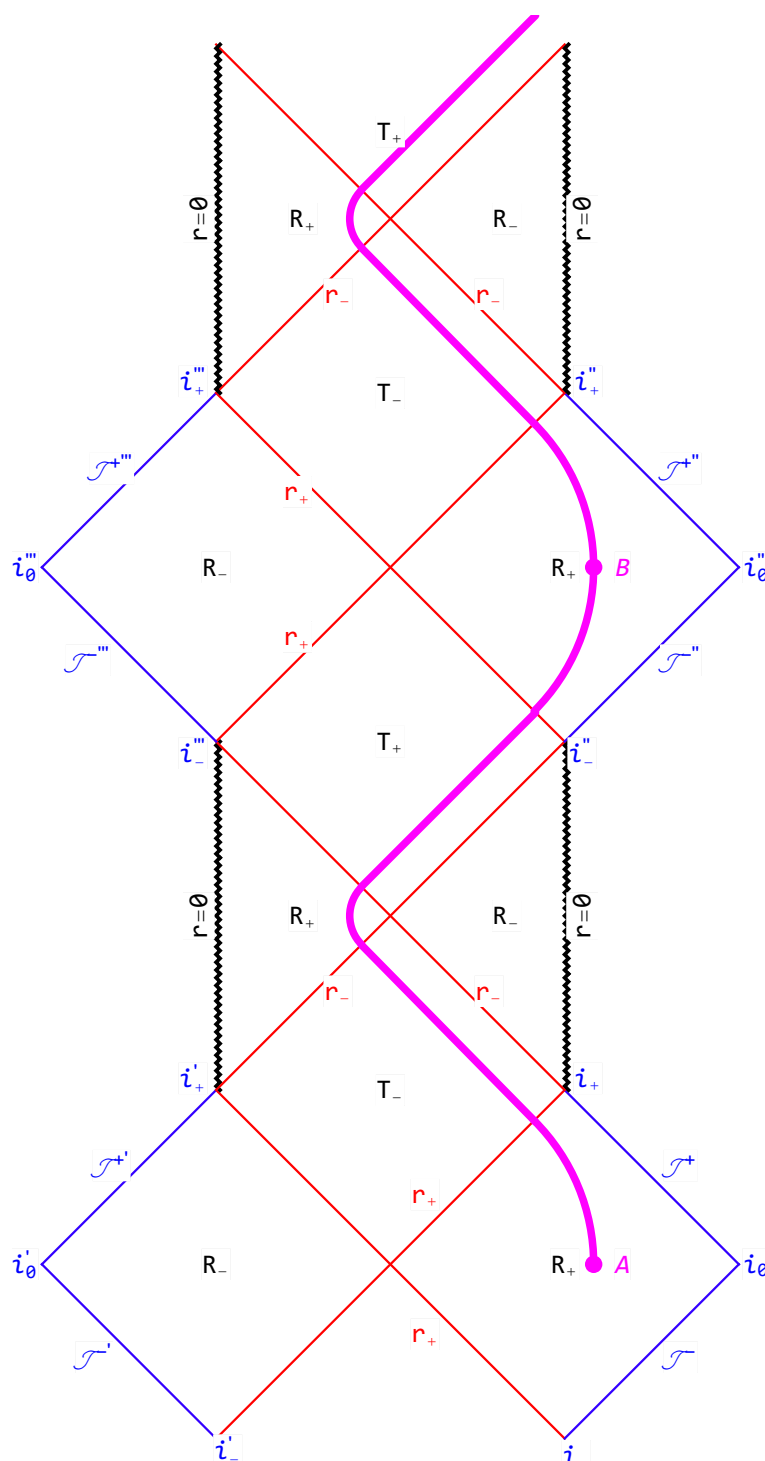


Figure 1. Carter-Penrose diagram for the spherically symmetric Reissner–Nordström black hole with electric charge $q = 0.99$. The spaceship starts from the point A at R_+ -region toward its multi-planetary future inside the black hole. The astronauts are planning to use the Einstein-Rosen bridge (magenta curve) and intersect both the black hole event horizon r_+ and Cauchy horizon r_+ at finite their proper time. After appearing near the black hole singularity at $r = 0$, the space ship uses its powerful engines to change the direction of motion and escape the tidal destruction at small radii. In result, the voyage is happily finishing at point B (may be at the Earth-like planet) in other infinite universe. The symmetry is in possibility to repeat the complete route of this voyage starting from the point B but only in the forward direction in time toward the another multi-planetary future. It is impossible to return back to the native Earth due to impossibility of any motion beyond the light cone. This is asymmetry of motion on the Einstein-Rosen bridge inside black hole.

2D presentation of the voyage through the interiors of Reissner–Nordstöm black hole with interiors by using the Einstein–Rosen bridge is shown at Figure 2. The electric charge of the black hole is $e = 0.99$. A test planet (or spaceship) with the electric charge $\varepsilon = -1.5$ is periodically orbiting around black holes with orbital parameters parameters $\gamma = 0.5$, $\lambda = 0.5$, corresponding to the maximal radius (apogee) $r_{\max} = 1.65$ and minimal radius (perigee) $r_{\min} = 0.29$, respectively, in dimensional units. The periodic planet geodesic trajectories (magenta curves, both at Figures 2 and 3), were calculated numerically with using equations of motion (3)–(6) for massive test particles ($\mu \neq 0$). Note, that the periodical motion of the test planet are limited in time due to energy losses in inevitable emission of the gravitational waves.

The picture at Figure 2 is geometrically absolutely *symmetric* or, in other words, it is completely or nicely *symmetric*. The geodesic trajectories of test planet ($\mu \neq 0$) at this Figure 2 and at Figure 4 (the red-colored curves), are numerically calculated [13–23], by using the corresponding equations of motion in the Kerr–Newman metric (3)–(6).

At the same time, this picture is misleading and physically controversial: Indeed the voyage is starting at apogee r_a from the position at point A, then reach the perigee r_p and return back the apogee r_a at the point B for a finite proper time, demonstrating the absolute geometric *symmetry*. Meanwhile, there is a crucial hitch: this apogee r_a at the point B is not in the native universe, but in the other quite distant universe, as it is clearly viewed at the Carter–Penrouse diagram at Figure 1. This hitch again destroys the Einstein–Rosen bridge *symmetry*.

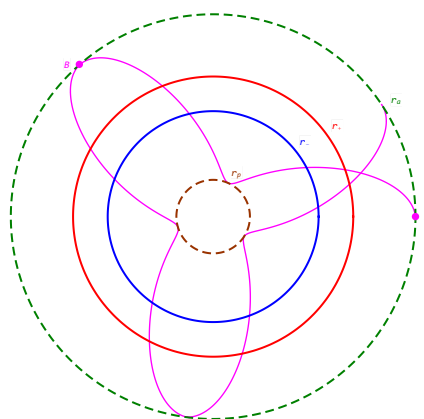


Figure 2. 2D presentation of the voyage through the Reissner–Nordstöm black hole interiors by using the Einstein–Rosen bridge. This picture is geometrically absolutely *symmetric* or, in other words, it is nicely *symmetric*. At the same time, this picture is misleading and physically controversial: Indeed the voyage is starting at apogee r_a from the position at point A, then reach the perigee r_p and return back the apogee r_a at the point B for a finite proper time, demonstrating the absolute geometric *symmetry*. Meanwhile, there is a crucial hitch: this apogee r_a at the point B is not in the native universe, but in the other quite distant universe, as it is clearly viewed at the Carter–Penrouse diagram at Figure 1. The apogee r_a and perigee r_p radii are shown by dashed circles. Respectively, the event radii of event horizon r_+ and Cauchy horizon r_- are shown by solid circles. The magenta curve here and at the Figure 3 is numerically calculated geodesic trajectory with using equations of motion (3)–(6) for massive test particles ($\mu \neq 0$).

Figure 4 shows the embedding diagram for the voyage through black hole interiors by using the Einstein–Rosen bridge. The embedding diagram is very useful for the training of intuitive understanding of the peculiarities of the enigmatic black holes. In this embedding diagram the Einstein–Rosen bridge connects two asymptotically flat universes like wormhole tunnel [24,25], but with the only one-way motion from the initial point A to the final point B. The geometrical *symmetry* of this embedding diagram is deceptive. In fact, this embedding diagram demonstrate the *asymmetric* space-time origin of the one-way Einstein–Rosen bridge (remember about loss-cone).

The completely *symmetric* picture of the periodic orbital motion of the test planet or spaceship around the central singularity of the Reissner–Nordstöm black hole inside the Cauchy horizon r_- is shown at Figure 3. The electric charge of the black hole is $e = 0.99$ and test planet (or spaceship) with the electric charge $\varepsilon = -1.5$ is periodically orbiting around black holes with orbital parameters $\gamma = 1.7$, $\lambda = 0.1$, corresponding to the maximal radius (apogee) $r_{\max} = 0.75$ and minimal radius (perigee) $r_{\min} = 0.09$, respectively. The *asymmetric* Reissner–Nordstöm bridge is only needed for penetration into this very exotic region at $0 < r < r_-$, where exist the nearly stable periodic orbits for test particles [26–32], which are very similar to the periodic orbits outside the black hole event horizon r_+ .

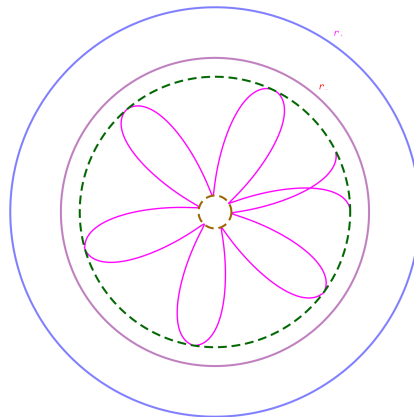


Figure 3. Both the geometrical and physical completely *symmetric* picture of the periodic orbital motion of the test planet or spaceship around the central singularity of the Reissner–Nordstöm black hole inside the Cauchy horizon r_- . The *asymmetric* Reissner–Nordstöm bridge is only needed for penetration into this very exotic region at $0 < r < r_-$, where exist the nearly stable periodic orbits for test particles, which are very similar to the periodic orbits outside the black hole event horizon r_+ . The apogee r_a and perigee r_p radii are shown by dashed circles. Respectively, the event radii of event horizon r_+ and Cauchy horizon r_- are shown by solid circles.

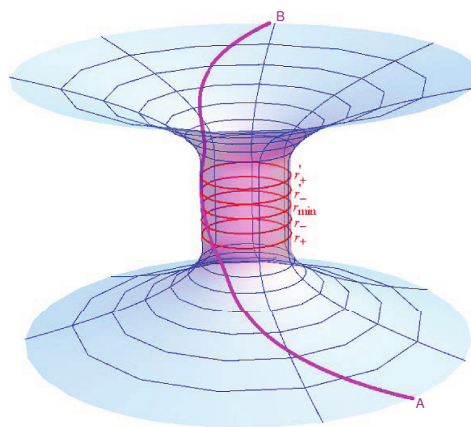


Figure 4. Embedding diagram for the voyage through black hole interiors by using the Einstein–Rosen bridge. This bridge connects two asymptotically flat universes like wormhole tunnel, but with the only one-way motion from the initial point A to the final point B . The geometrical *symmetry* of this embedding diagram is deceptive. In fact, this embedding diagram demonstrate the *asymmetric* space-time origin of the one-way Einstein–Rosen bridge (remember about loss-cone).

4. Symmetry and Asymmetry of Test Particle Trajectories Near Rotating Black Hole

Figures 5–7 demonstrate both *symmetry* and *asymmetry* features of massive and massless particle trajectories plunging into rotating Kerr black hole with spin $a = 0.9982$. **Magenta** arrows shows the direction of the black hole rotation in accordance with the gimlet rule. The **multi-colored** curves at Figures 7 and 8 are the geodesic trajectories for massless test particles like photons ($\mu = 0$) numerically calculated with using equations of motion (3)–(6). By approaching the black hole the trajectories of all particles, both massive and massless ones, are infinitely winding up on the black hole event horizon in the direction of the black hole rotation and at the fixed latitudes. This winding up is a manifestation of *symmetry* of behaviour of all trajectories, plunging into rotating black hole. At the same the direction of the black hole rotation is a corresponding manifestation of *asymmetry* of the gravitational field of the Kerr metric.

At last, for completeness of black hole *symmetric* and *asymmetric* properties, at Figure 8 is shown the trajectory of the test planet ($\mu \neq 0$) with parameters $\gamma = 1$, $\lambda = 1$ and $Q = 0.5$. This test planet is plunging into the spherically *symmetric* and nonrotating Schwarzschild black hole (with both the spin $a = 0$ and electric charge $q = 0$), starting from the radial distance $r = 6$. It must be especially checked that the traversable (though only one-way in time and direction) Einstein-Rosen bridge is absent at all inside the Schwarzschild black hole (see for details, e.g., [4,6]).

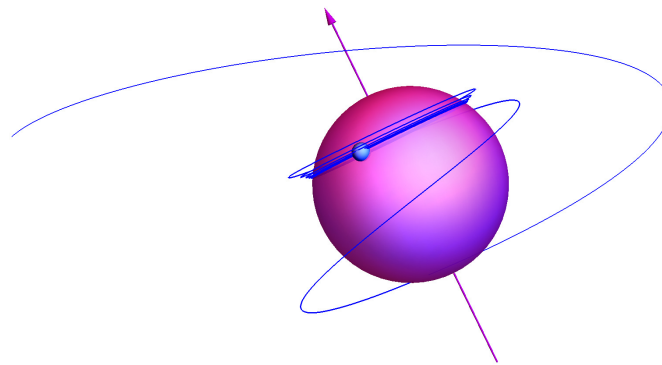


Figure 5. Trajectory of the test planet with parameters $\gamma = 0.85$, $\lambda = 1.7$ and $Q = 1$, which is plunging into the fast rotating Kerr black hole with spin $a = 0.9982$. This planet is winding up on the black hole event horizon higher the equatorial plane. The **blue** curve here and in the Figure 6 is the numerically calculated geodesic trajectory with using equations of motion (3)–(6) for massless test particles like photons ($\mu = 0$).

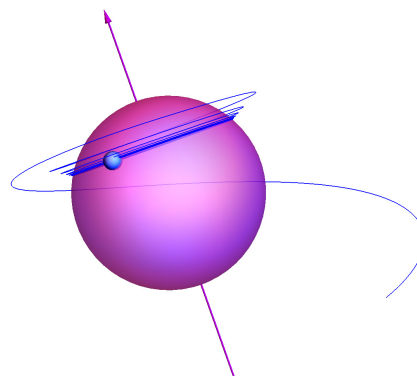


Figure 6. Trajectory of the test planet with parameters $\gamma = 0.85$, $\lambda = 1.7$ and $Q = 1$, which is plunging into the fast rotating Kerr black hole with spin $a = 0.9982$. This planet is winding up on the black hole event horizon higher the equatorial plane. The **blue** curve here and in the Figure 7 is the numerically calculated geodesic trajectory with using equations of motion (3)–(6) for massless test particles like photons ($\mu = 0$).

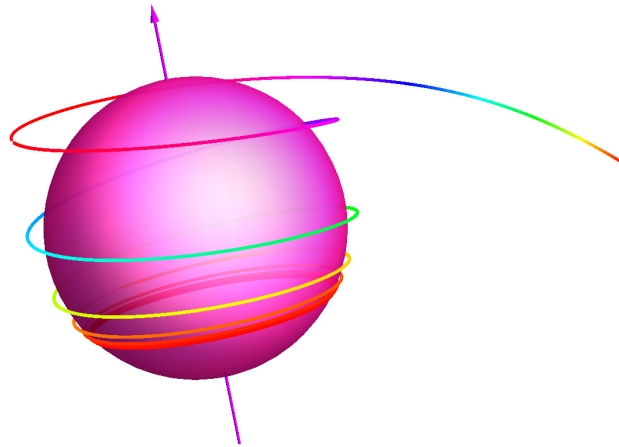


Figure 7. Trajectory of the photon with parameters $\lambda = 2$ and $Q = 1$, which is plunging into the fast rotating Kerr black hole with $a = 0.9982$ and is winding up on the black hole event horizon below the equatorial plane. The multi-colored curve is the numerically calculated geodesic trajectory with using equations of motion (3)–(6) for massless test particles like photons ($\mu = 0$).

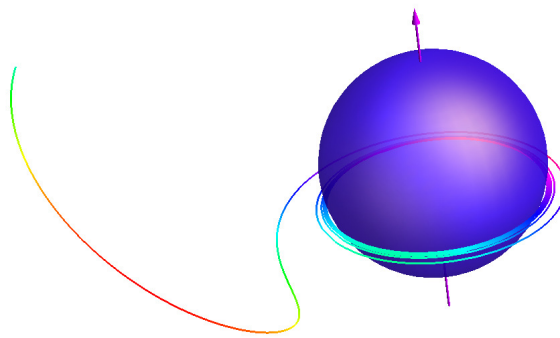


Figure 8. Trajectory of the photon with parameters $\lambda = -6.5$ and $Q = 4$, which is plunging into the fast rotating Kerr black hole with $a = 0.9982$ and is winding up on the black hole event horizon below the equatorial plane.

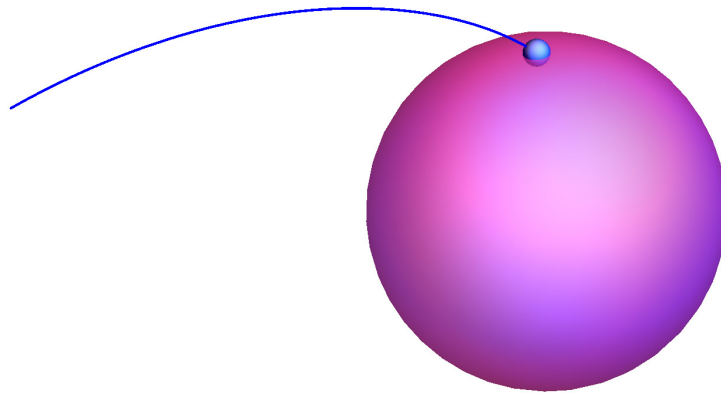


Figure 9. A trivial but though very expressive trajectory of a test planet ($\mu \neq 0$) with parameters $\gamma = 1$, $\lambda = 1$ and $Q = 0.5$, which is plunging into the the spherically *symmetric* and nonrotating Schwarzschild black hole (with both the spin $a = 0$ and electric charge $q = 0$). The starting point for this crazy voyage is at the radial distance $r = 6$.

5. Conclusions and Discussion

It is demonstrated the symmetry and asymmetry of the voyage on one-way Einstein-Rosen bridge inside black hole toward the endless multiplanetary future. The apparent symmetry of both the Carter-Penrose and embedding diagrams is mainly related with a pure geometrical vision of this phenomenon. Quite the contrary, the physical (space-time) vision elucidates the absolute asymmetry of the Einstein-Rosen bridge due to existence of the light-cone limitation for possible motions.

Note, that the traversable (though only one-way in time and direction) Einstein-Rosen bridge exist only in the case of both rotating Kerr $a \neq 0$ and electrically charged Reissner–Nordström $q \neq 0$ black holes. It is absent at all inside the Schwarzschild black hole (see for details, e.g., [4,6])

The infinite winding up of trajectories of all particles on the black hole event horizon is a manifestation of *symmetry* of behaviour of all trajectories, plunging into rotating black hole. At the same the direction of the black hole rotation is a corresponding manifestation of *asymmetry* of the gravitational field of the Kerr metric.

Author Contributions: Investigation, V.I. Dokuchaev and K.E. Prokopiev. The authors contributed to the paper equally.

Acknowledgments: We are grateful to E. O. Babichev, V. A. Berezin, Yu. N. Eroshenko and A. L. Smirnov for stimulating discussions.

Conflicts of Interest: The authors declare no conflict of interest.

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