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Article

Characterizing Low-Energy Charged Particles in the Magnetosphere with the LEM CubeSat Spectrometer Project: Detector Concept and Hardware Characterisation

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Abstract: An accurate flux measurement of low-energy charged particles trapped in the magnetosphere is necessary for space weather characterization and to study the coupling between the lithosphere and magnetosphere, which allows for the investigation of the correlations between seismic events and particle precipitation from Van Allen belts. In this work, the project of a CubeSat space spectrometer, the Low-Energy Module (LEM), is shown. The detector will be able to perform an event-based measurement of the energy, arrival direction, and composition of low-energy charged particles down to 0.1 MeV. Moreover, thanks to a CdZnTe mini-calorimeter, the LEM spectrometer also allows for photon detection in the sub-MeV range, joining the quest for the investigation of the nature of Gamma-ray bursts. The particle identification of the LEM relies on the $\Delta E - E$ technique performed by thin silicon detectors. This multipurpose spectrometer will fit within a $10 \times 10 \times 10$ cm³ CubeSat frame, and it will be constructed as a joint project between the University of Trento, FBK, and INFN-TIFPA. To fulfil the size and mass requirements, an innovative approach, based on active particle collimation, was designed for the LEM; this avoids the heavy/bulky passive collimators of previous space detectors. In this paper, we will present the LEM geometry, its detection concept, the results from the developed GEANT4 simulation, and some characterisations of a candidate silicon detector for the instrument payload.

Keywords: low-energy module; low-energy particles; gamma-ray bursts; space weather; cubesat; $\Delta E - E$ technique

1. Introduction

The Low-Energy Module (LEM) will be a compact spectrometer able to perform an event-based measurement of the energy, direction, and composition of low-energy charged particles, in particular, down to 0.1 MeV for electrons. The physics goal of this detector is the monitoring of the magnetosphere and ionosphere environment. It is known that the measurements of the fluxes of low energetic particles may allow the characterisation of the coupling between the lithosphere, atmosphere, ionosphere, and magnetosphere. In particular, earthquakes are dynamic processes caused by continuous and

slow strain accumulation. From studies on fault rupture mechanics, seismic wave propagation, and geophysical parameters measured in the ionosphere and the low magnetosphere, some anomalies correlated with catastrophic events were discovered. Moreover, statistical evidence of a temporal correlation between particle precipitations from Van Allen belts and strong seismic events has been pointed out [1]. These observations motivate interest in further detailed measurements of electron fluxes in the energy window 0.1–7 MeV, which may be a promising channel for identifying possible seismic precursors. Another interesting case study for the LEM instrument is its application to space weather. Severe space weather storms can cause power outages and telecommunication alterations. For this reason, the construction of new instruments to monitor and (possibly) predict the effects of solar activity on Earth is crucial.

The LEM will be a particle telescope performing an event-based measurement of the energy, direction, and composition of low-energy charged particles, in particular, electrons down to 0.1 MeV. This capability is not possible with the existing detectors, for which the possibility of an event-based PID or the possibility of monitoring the particle flux from different directions at the same time fails or are not able to measure the directions of low-energy particles because of the multiple scattering occurring in the first layer of a particle-tracking configuration.

2. Investigation of the magnetosphere: Correlation between Particle Precipitations and seismic events

The magnetosphere [2] is the complex environment surrounding Earth's atmosphere in which the magnetic field generated by Earth's core is dominant. Earth's fused core, with its continuously flowing currents, generates a relatively weak magnetic field (of about 6×10^{-5} T at Earth's surface near the poles) well described by a dipolar field. However, at several Earth's radii, the magnetic dipole field is strongly deformed by the presence of the solar wind, a plasma of electrons and ions moving outward from the sun. Van Allen radiation belts inside the magnetosphere contain energetic ions and electrons that experience long-term magnetic trapping. However, near magnetopause, particle trapping is prevented since the magnetic field is not stable in time. Therefore, radiation belts are located below about 7 Earth's radii at the equator. On the other hand, at low altitudes, the atmosphere prevents particle trapping. The main reason is that charged particles, like protons or electrons, lose energy during collisions. Consequently, the region in which radiation belts are located is above 200-1000 km. Particles' dynamics in the radiation belts is characterised by a superposition of three different motions acting on different time scales. In Figure 1, a GEANT4 [3] simulation of a 100 MeV trapped proton inside the dipolar approximation of Earth's magnetic field (with the magnetic moment pointing upwards). A charged particle will rotate in a dipolar field following a helicoidal trajectory with a high rotation frequency. On slower time scales, the particle travels along the field lines, bouncing between the two mirror points. Again, on slower time scales, the particles, while revolving about the field lines and bouncing, also follow a longitudinal drift motion. Proton energies extend up to several hundred MeV even in the LEO. This aspect makes protons one of the most critical hazards for satellites at these altitudes. Energetic electrons are the other most abundant population of particles lying in the radiation belts. Their energy extends up to several MeV, and their average fluxes experience time variations related to geomagnetic activities and the succession of solar cycles. There are several processes that can cause a precipitation of the trapped charged particles from the radiation belts. These processes are not fully understood yet, but the primary cause is thought to be electromagnetic fluctuations in the radiation belt. These electromagnetic fluctuations can be produced through solar-magnetic storms, lightning storms, man-made electromagnetic emissions and seismic activity.

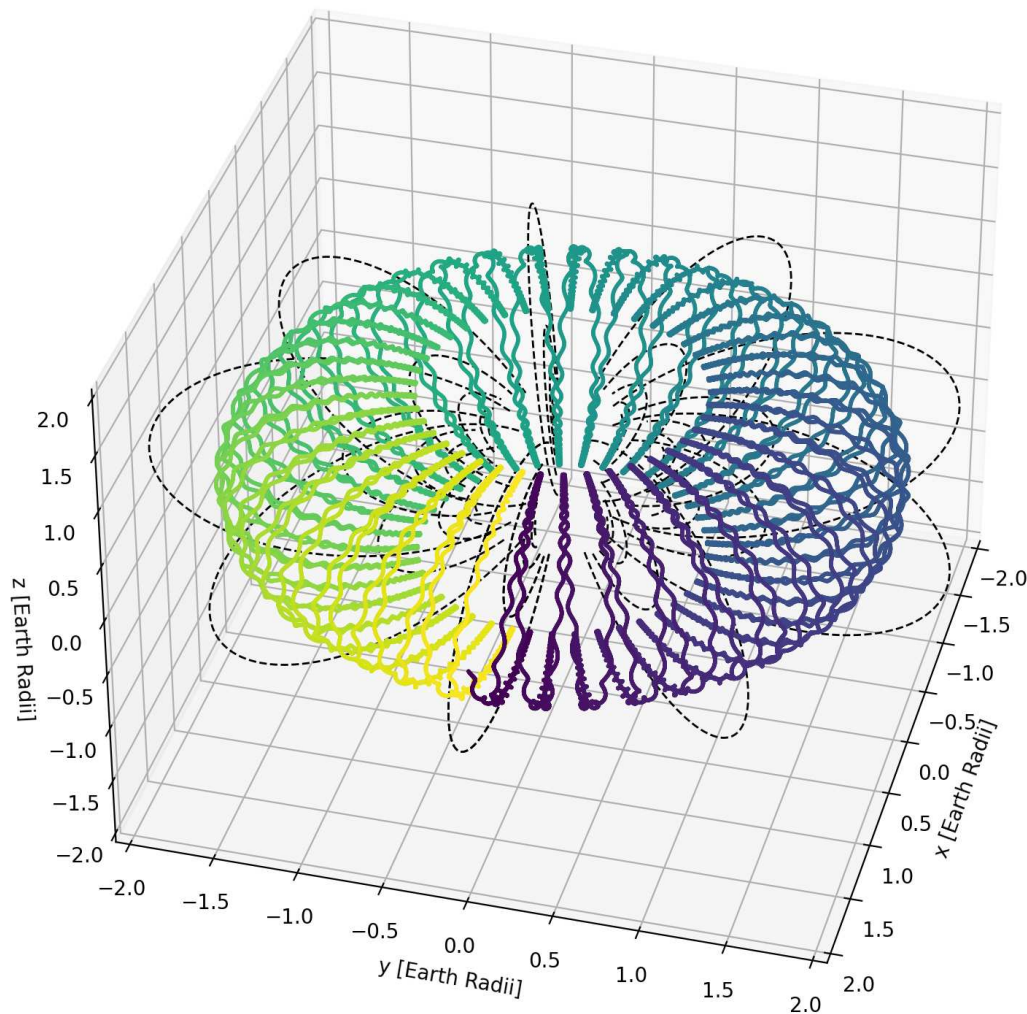


Figure 1. GEANT4 simulation of the trajectory of a 100 MeV proton trapped inside a dipolar approximation of Earth's magnetic field. The magnetic moment is pointing upwards. The drift motion of the proton goes counter clockwise looking from the top. Dashed black lines depict some magnetic field lines. The colour scale, from purple to yellow, encodes the time of the propagation of the particle.

There are many literature models trying to describe the coupling between the Lithosphere-Atmosphere-Ionosphere-Magnetosphere (LAIM). An extensive review can be found in reference [4]. However, since the lithosphere is a very complex and heterogeneous system, there is the need to find a very general model avoiding the introduction of assumptions which are valid only in some specific cases. For monitoring seismic phenomena, in particular, space is a privileged point from which it is possible to measure many observable quantities. Indeed, satellites and satellite constellations provide remote sensing of variables within the atmosphere, ionosphere, and magnetosphere. Improving data quality will allow a better comprehension of such systems, particularly their motions, energy exchanges, and dynamics.

Monitoring of trapped particle precipitations is a channel to study the LAIM coupling. Indeed, among the many LAIM coupling models it is interesting to mention the Magnetosphere-Ionosphere-Lithosphere Coupling (MILC) model, described in [5]. Some observations corroborate the presence of co-seismic perturbations of the magnetosphere [5,6]. Observations of co-seismic alterations of the magnetosphere environment and consequent observation of particle precipitations are crucial points for validating such models.

3. The Current Landscape of Space-Based Particle Detectors

Extensive literature exists about particle detectors in space using silicon technologies for ΔE - E measurements. Here, a not exhaustive list of existing detectors will be discussed as an example. The instruments under examination are: the Instrument for the Detection of Particles (IDP) on the DEMETER microsatellite [7–9], the High-Energy Particle Package (HEPP-H and HEPP-L) on CSES [10–13], the Mars Energetic Particle Analyzer (Mars-EPA) on the Tianwen-1 mission [14–16], and the Radiation Assessment Detector on the Curiosity rover [17–21].

Some of the most important features of these instruments are listed in Table 1. Even though all of these experiments have different scientific purposes and goals, their detection concepts and schemes are very similar, allowing a comparison between their structure, size, components, and performances. By comparing the six detectors studied, we can conclude that the larger the number of layers inserted into the design, the better the performances in detecting energetic particles and particle identification. On the other hand, to minimise the low-energy threshold, one has to minimise the thickness of the ΔE layer. For instance, the Mars-EPA, can detect electrons in the energy range of 0.1–2 MeV by adopting a ΔE layer made of Passivated Implanted Planar Silicon (PIPS) detectors with a thickness of 15 μm . Finally, the use of an inorganic scintillator as a calorimeter could be problematic. In particular, many scintillator crystals, such as Sodium-Iodide or Caesium-Iodide, are very fragile and hygroscopic. These aspects will unavoidably result in the introduction of mechanical supports or metallic wrapping, providing additional dead layers in which particles could deposit part of their energy.

Furthermore, it is required that the LEM is compact (within $10 \times 10 \times 10 \text{ cm}^3$) and that it can monitor the particle flux in a large field of view from different directions at the same time. These capabilities are not simultaneously fulfilled by the past detectors. Therefore, a different and innovative design is required for the LEM.

Table 1. Summary of some features of the detectors studied in this section. The references from which I extracted the information are quoted within the text.

Instrument	Size Weight	Directions	Angular Resolution	Energy Range	PID	Detector Elements
IDP DEMETER	525 g	1	FOV 32 deg.	e: [0.07, 0.8] MeV	No	Silicon Diode
RAD Curiosity	$\sim 10 \times 10 \times 10 \text{ cm}^3$	Complex segmentation	FOV 36.7 deg.	e: [0.1, 20] MeV p: [5, 200] MeV α : [5, 200] MeV I.Z: [10, 300] MeV	Yes	PIPS (3 segments) CsI(Tl) Plast. Scint.
HEPP-L	Large Collimators	5 Narrow 4 Wide	FOV 6.5 deg. FOV 15 deg.	e: [0.1, 3] MeV p: [2, 20] MeV	Yes	Si det. (2 layers) Plast. Scint.
Mars-EPA	$270 \times 180 \times 148 \text{ cm}^3$	1	FOV 60 deg.	e: [0.1, 12] MeV p: [2, 100] MeV α : [25, 400] MeV I.Z: [25, 400] MeV	Yes	PIPS (2 layers) CsI(Tl)

4. The LEM Concept: The Active Collimation Technique

The idea allowing for a reduction in the weight and size of the LEM detector relies on the active collimation technique. More precisely, a drilled plastic scintillator is acting as a veto. Only particles with the directions aligned with 1 of the 16 channels are detected by 1 of the 16 silicon sensor pairs. Thus, the direction information is obtained. Particles with the different/unknown directions are stopped in the aluminium shield or will release a signal in the drilled plastic scintillator veto. This technique is an alternative to the tracking one affected by the multiple scattering problem. On the other hand, the low density of the plastic scintillator veto avoids the significant weight required by a totally passive metallic collimator. However, the price to pay is a relatively high veto rate. This high

veto rate will unavoidably result in an enhancement of the dead time of the detector. For this reason, a small drilled aluminium shield is still necessary to suppress very low energetic particles.

In Figure 2, the detection concept and a schematised cross-section of the instrument are shown. In Figure 3, the LEM geometry is displayed within the developed GEANT4 simulation. From the top, we can see the drilled aluminium mask, in transparent grey, suppressing the flux of very low energetic particles. Below the aluminium shield, the active anti-coincidence is obtained by using a drilled plastic scintillator, displayed in transparent blue, (polyvinyl toluene). The aluminium drilled mask and the drilled anti-coincidence detector (ACD) define the so-called active collimator.

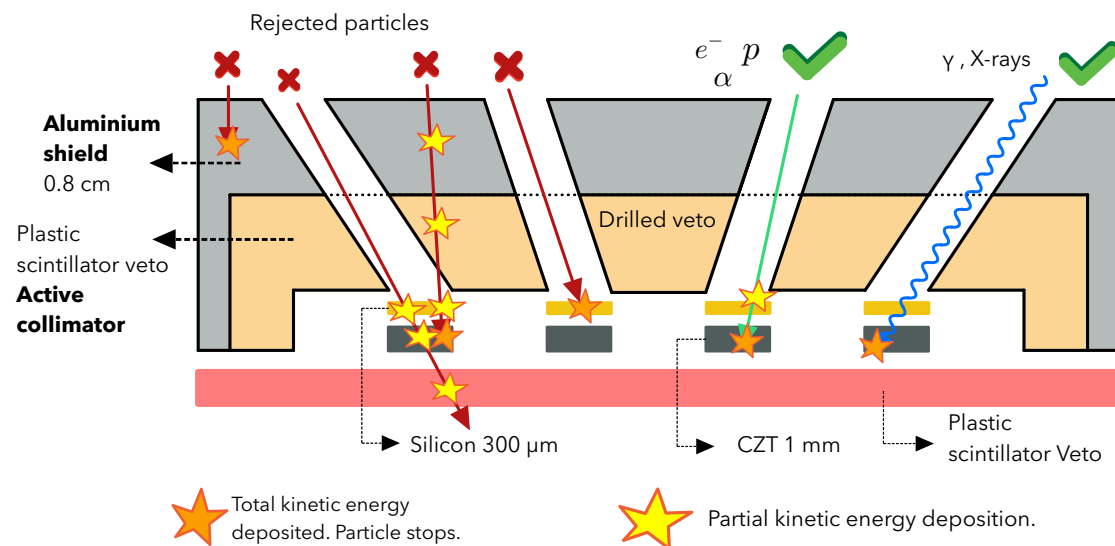


Figure 2. Detection concept embedded within the detector's geometry. In the picture, red trajectories represent discarded events, green trajectories represent good/accepted particle events, the blue trajectory represents a good/accepted photon event.

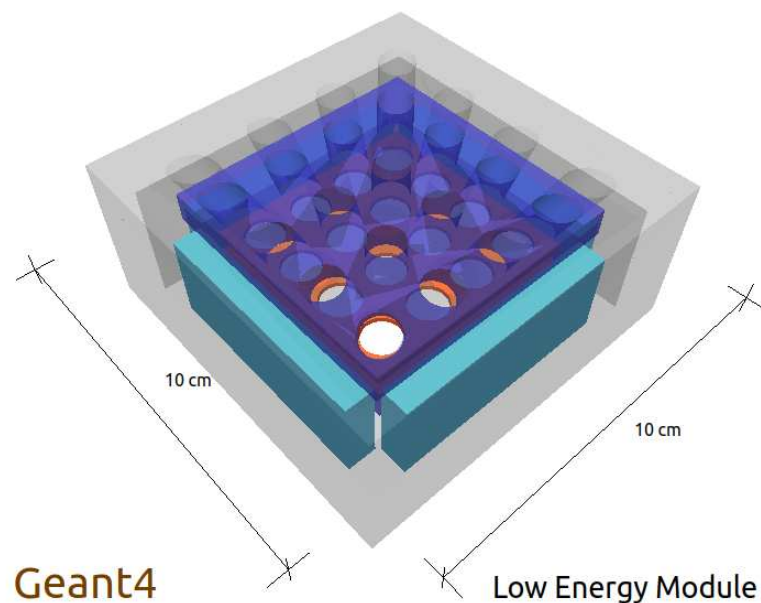


Figure 3. Visualization of the Low Energy Module (LEM) geometry developed with the GEANT4 framework [3].

For an LEM operating in Low Earth Orbit (LEO), an aluminium thickness larger than 0.5 cm is necessary to reduce the veto rate from several MHz to the affordable rate of \sim kHz. Below the active collimation system, we place the 16 independent $\Delta E - E$ modules. These $\Delta E - E$ modules will measure the angular flux of particles crossing the veto channels ($\varnothing 1 \text{ cm} \times 1.3 \text{ cm}$), determining one specific solid angle in the sky with a resolution of about 7° . The sizes of the commercially available PIPS detectors ($50 \text{ mm}^2 - \varnothing 8 \text{ mm}$ each) have been considered to define a realistic geometry in the detector simulation. The ΔE detector consists of a $100 \mu\text{m}$ thick PIPS detector while the E detector is a CdZnTe (or CZT) detector with a 1 mm thickness. These two $\Delta E - E$ layers allow a good particle identification in the energy ranges of approximately 0.1–10 MeV for electrons, 3–30 MeV for protons, and 10–100 MeV for alpha particles. A bottom plastic scintillator (ACD), not shown in Figure 3, is added at the very end of the LEM to ensure that the energy release is confined within the above layers. Plastic scintillator (ACD), displayed in cyan, is inserted on all four sides to ensure lateral confinement. In particular, particle identification (PID) is not possible for the energetic particles crossing the ACD nor for slow particles stopped in the front PIPS. Events with an undefined direction are rejected thanks to a signal released in the active veto/collimator. Finally, events that are fully contained within the LEM, are selected. In this very last case, the direction is well defined, and it is also possible to perform an accurate PID. Thanks to the high density and high averaged atomic number of CZT [22,23], the LEM can identify low-energy γ -rays converting in the CdZnTe (CZT) mini-calorimeter and using all the surrounding low-Z sensors as anticoincidence. The ability to observe energetic photons will allow for the additional use of this compact particle spectrometer as a Gamma-ray burst (GRB) monitor [24].

5. Performance Characterisation with GEANT 4 Simulation

5.1. Expected performance for low energy charged particles

The detection concept adopted in the LEM is a consolidated technique denominated $\Delta E - E$ [25–27]. Basically, a $\Delta E - E$ particle spectrometer is composed of a thin detection layer and a thicker one behind. When a particle impinges on the spectrometer, if the kinetic energy is enough, the particle can cross the first layer, releasing a part of its kinetic energy ΔE . Then, the residual kinetic energy E can be deposited entirely within a second, thicker layer. This experimental layout allows for particle identification by measuring the energy deposited in the thinner layer, the ΔE energy, as well as the energy deposited in the thick sensor, the E energy. If a non-relativistic particle passes through a thin detector layer, the energy deposited, ΔE , will be velocity dependent:

$$\Delta E \approx \frac{Z^2}{\beta^2} \quad (1)$$

where Z is the projectile's charge, and β is its velocity in natural units. On the other hand, the residual kinetic energy, E , of a non-relativistic particle stopping in a subsequent thick detector is also velocity dependent:

$$E = mc^2(\gamma - 1) \approx \frac{1}{2}m(\beta c)^2 \quad (2)$$

Therefore, in a $\Delta E - E$ spectrometer, a useful PID classifier can be defined in the following way:

$$\text{PID}_{\text{classifier}} = \log_{10} \left[\frac{\Delta E}{1 \text{ MeV}} \frac{E}{1 \text{ MeV}} \right] \approx \text{constant} + \log_{10} Z^2 \left(\frac{mc^2}{1 \text{ MeV}} \right) \quad (3)$$

Thus, for non-relativistic particles, this $\text{PID}_{\text{classifier}}$ is mainly dependent only on the particle's mass and charge. For the LEM detector, this approximation is very good for protons, alpha, and other nuclei, but, for electrons in the LEM kinetic energy range, the non-relativistic approximation fails, and the PID classifier for electrons will grow roughly according to $\log_{10} \frac{E}{1 \text{ MeV}}$. However, a good identification of the electrons from the protons, based on this classifier, is still achieved thanks to the fact that the

proton mass is 2000 times larger than the electron mass. In Figure 4, the results from a GEANT4 [3] simulation are shown. In particular, the PID vs. the energy identification capability for the case of a mini-calorimeter made of 500 μm of silicon (left plot) is compared with the case of a mini-calorimeter composed of a 1 mm thick CZT.

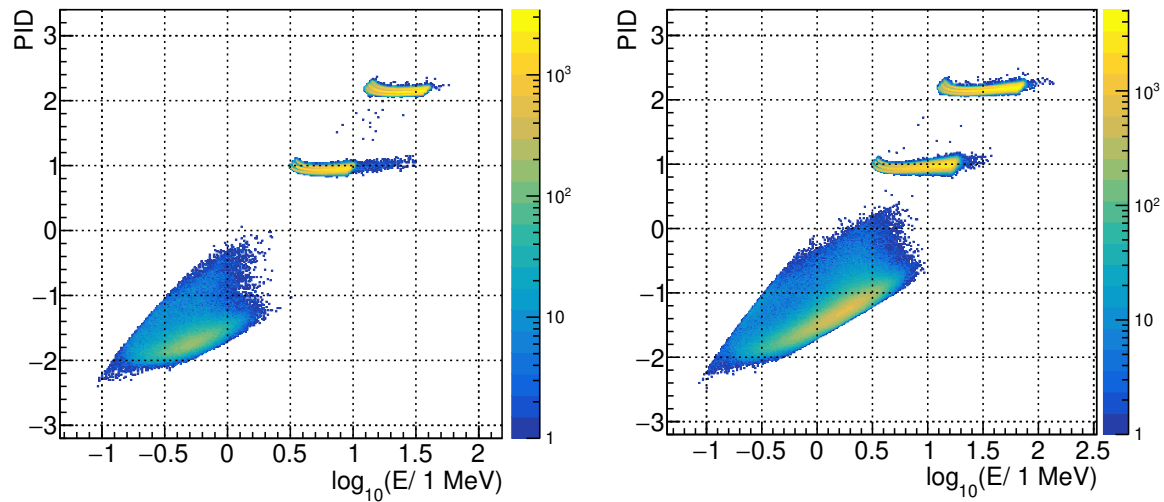


Figure 4. PID classifier vs. kinetic energy. **Left:** 100 μm –500 μm PIPS detectors. **Right:** 100 μm PIPS and 1 mm CZT detectors. The three different clusters in each plot represent (from the top to the bottom): alpha particles, protons, electrons.

Using the results from a long run of the simulation where particles are isotropically generated, it is possible to quantify the angular resolution of the detector. The results are reported in Figure 5. In the notation used, PIPS CZT detectors are lying in the xy plane. To represent on a two-dimensional plane the vectors lying on the unitary sphere, encoding the incident direction of the projectile particle, it has been adopted a polar projection on which $X = \theta \cos \phi$ while $Y = \theta \sin \phi$. The z direction, perpendicular to the plane of the graph, corresponds to the zenith direction. To understand which portion of the sky is subtending each collimator, events detected by adjacent PIPS-CZT pairs are visualised with different colours. Therefore, each cluster of particles corresponds to a class of events detected by the same $\Delta E - E$ module. For each channel the angular resolution for protons and alpha particles is $\sim 6^\circ$ whereas is $\sim 7.5^\circ$ for electrons. The worst angular resolution of electrons is mostly due to a larger effect of particle multiple scattering crossing the collimator edges. The overall FOV for the LEM is $\sim 60^\circ$.

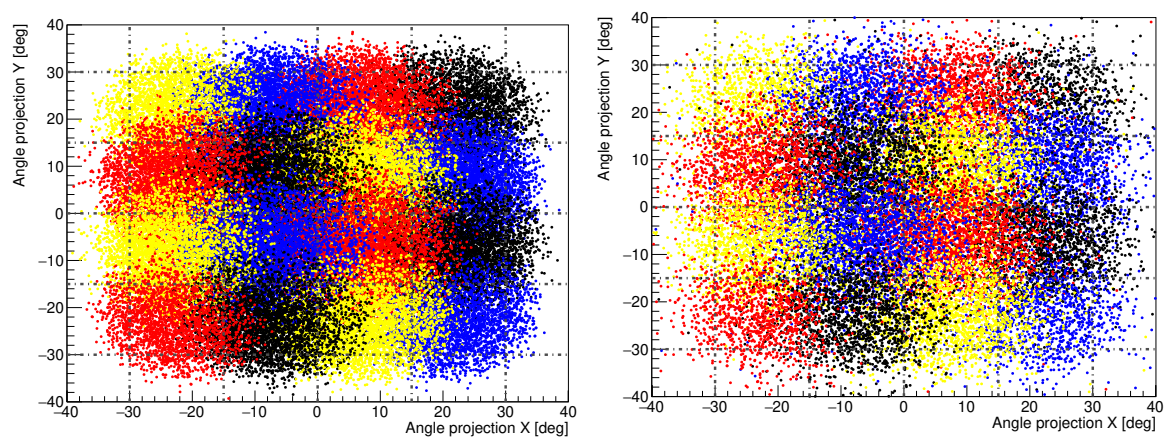


Figure 5. Angular projection of the projectile's incident direction. In this graph, the xy plane is parallel to the detectors. The colour code identifies the $\Delta E - E$ silicon sensor pair that detects the particle. Left plot is relative to proton and alpha particles whereas in the right plot electron events are shown.

5.2. Expected performance for gamma rays

Gamma ray Bursts (GRBs) are one of the most extreme and violent phenomena in the universe. GRBs [24], by definition, are short and very intense bursts of gamma rays and lower-energies photons sometimes accompanied by optical afterglows. GRBs represent the most luminous object in the sky reaching luminosity of $10^{51} - 10^{52}$ ergs s^{-1} . They offer the possibility to explore the early universe, study star formation, validating fundamental theories and principles [28–30]. Studies on the hardness-duration (T_{90}) of the detected GRBs by several space-based instruments determined the presence of two populations of GRB: long ($T_{90} > 2$ s), associated with super novae explosions, and short ($T_{90} < 2$ s), associated with binary neutron star mergers. Nevertheless, there are hints for additional populations/sub-classes of GRBs [31]. Increasing the number of available GRB monitors will allow to cover a bigger portion of the sky allowing better statistics, better coverage, and improving the multimessenger astronomy capabilities such as in the case of the GW170817 observation [32]. Nevertheless, up to now, no coincidences between GRBs and neutrino events have been observed [33]. Distributed architectures for GRB monitoring [34] will allow an improvement limits on the neutrino fluxes emitted from GRBs.

Recently, GRB20221009A, the strongest GRB event observed until now, was also detected by particle detectors such as HEPP-L onboard the CSES-01 satellite [35], and by electron and proton detectors on board the four spacecraft of the NASA THEMIS mission [36]. In both cases, the GRB was detected due to the secondary production in the interaction between soft gamma rays with the passive materials of the collimators.

Even though GRB221009-like events could be rare, LEM detector would be able to catch GRBs both monitoring charged particle fluxes and with the standard calorimetric technique (with CdZnTe mini-calorimeters).

To quantify the advantage of the use of CZT sensors in the LEM also as a monitor for GRBs, a comparison of the relative photon detection efficiency for the two mini-calorimeter configurations is shown in Figure 6.

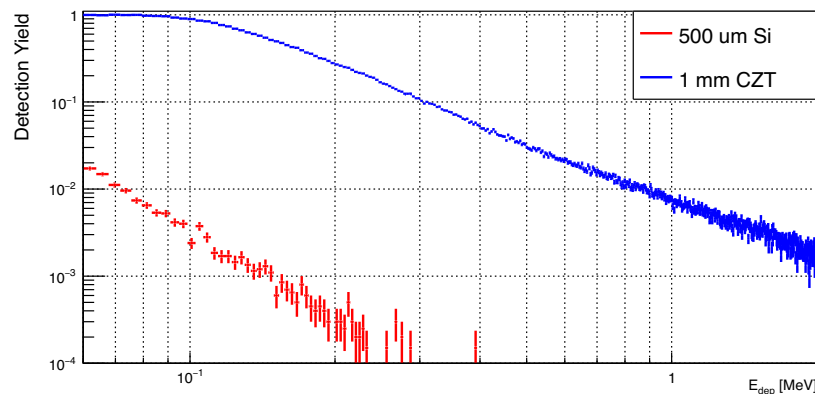


Figure 6. Improvement of detection yield for γ -rays thanks to CZT sensors. The legend refers to the material and the thickness considered for the E detector.

The use of CZT sensors is promising to monitor the sub-MeV part of the GRB spectrum with a very good energy resolution.

6. Characterisation of Fully Depleted Silicon Detector prototypes

This chapter will be devoted to the activities regarding the Hardware R&D, which is accomplished at INFN-TIFPA for the LEM project. Currently, we have carried out the characterisation of the Passivated Implanted Planar Silicon (PIPS) detector 500 μ m, AP-CAM25 manufactured by MIRION, with a built-in preamplifier. The goal of these measurements were the tuning of the LEM simulation and the test of the silicon response to detected particles having very different dE/dx .

From the electrical point of view, the working principle of a fully depleted silicon detector detector is basically the same as that of a reverse biased pn-junction. A particle entering and releasing some energy in the depletion region forms electron-hole pairs. In the case of silicon, the energy required for forming one pair is about $\varepsilon \simeq 3.6$ eV. As a result, the number of electron-hole pairs generated inside the depletion layer will be $N_{eh} = \eta E_{in} / \varepsilon$ where, η is the quantum efficiency, and E_{in} is the energy deposited within the detector depletion region. The thickness of the depletion region depends on the bias voltage applied. The higher the bias voltage, the more extensive the region over which a strong electric field will separate electron-hole pairs. Electrons and holes are then accelerated towards the cathode and the anode, where they are collected. However, the thickness is limited by the dimensions of the wafer. As a result, the depletion layer will remain constant above the depletion voltage.

6.1. Calibration with γ -ray sources

To calibrate the response of silicon detector in low energy range two *gamma*-ray external sources were used: ^{241}Am standard calibration source and a LYSO (Lutetium-Yttrium Oxyorthosilicate) crystal scintillator.

In particular, ^{176}Lu is one of the constituents of the LYSO crystal scintillator. ^{176}Lu is a naturally available (2.6%) and long lived (40 Gyr) isotope of Lutetium. It decays β^- to ^{176}Hf . The decay chain involves the emission of three γ rays with energy respectively 307 keV, 202 keV, and 88 keV [37]. In addition, there is a peak at ~ 55 keV due to the $K\alpha$ emission line. The expected activity of LYSO material is ~ 40 Bq/g. A thin LYSO scintillator (8 grams) was coupled to an Hamamatsu R5946 PMT in order to tag the ^{176}Lu β^- decay with 100ns coincidence with the silicon detector. In Figure 7 (left), it is possible to see the energy spectrum measured by the silicon detector detector exposed to the LYSO radioactivity. To verify the calibration curve at low energy, we used the 59.5 keV gamma line emitted by the ^{241}Am calibration source. Moreover, removing the source we also measured the "pedestal", which is the signal detected by the detector, due to the electronic noise, in absence of energy deposition. Figure 7 (Right) shows the results from the measurements with the Americium source. As it is possible to see, the pedestal measurement allows the characterisation of the electronic noise and background in the detector.

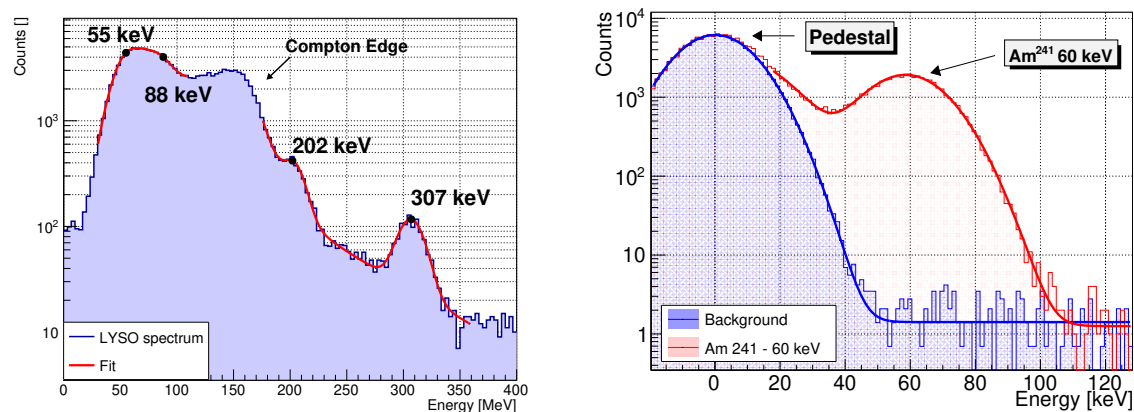


Figure 7. (Left) LYSO spectrum measured with the fully depleted silicon detector. (Right) Spectrum of the ^{241}Am radioactive isotope in the silicon detector.

Figure 8 shows the calibration curve and the detector's resolution estimation. On the left-hand side of the plot is reported the detector's resolution estimation. This calibration procedure provides a calibration constant: (36.7 ± 0.3) mV/MeV and an energy resolution of: $(11.1 \pm 0.4^{\text{stat}} \pm 1.2^{\text{syst}})$ keV.

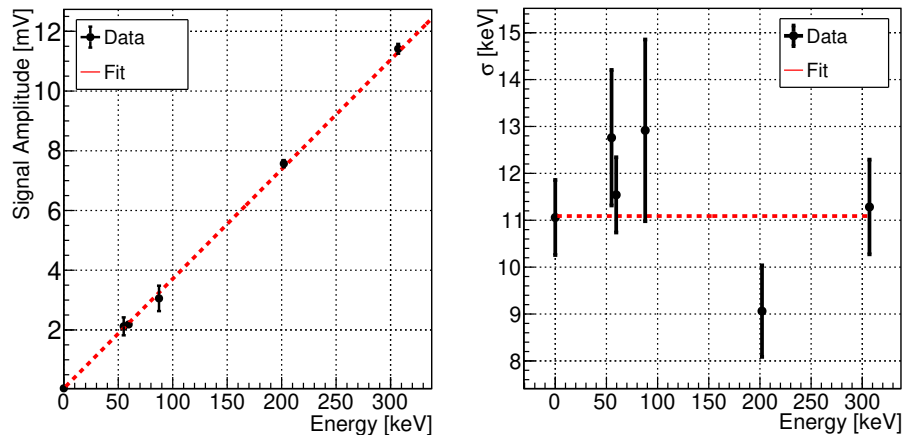


Figure 8. Calibration curve and measurement of the resolution of the silicon detector.

6.2. Calibration with cosmic muons (MIP)

To test the PIPS response to Minimum Ionising Particles (MIPs) the natural atmospheric muons ($\sim 100 \text{ Hz m}^{-2}$) [38] has been acquired asking for a coincidence of the PIPS detector with two Plastic Scintillators placed in a telescopic configuration.

In Figure 9 are reported the data from the calibration with muons and the results from the curve fitting procedure. The MPV was found to be $\text{MPV} = (5.626 \pm 0.016) \text{ mV}$, while the gaussian resolution was $\sigma = (0.53 \pm 0.03) \text{ mV}$. From the comparison with a MonteCarlo GEANT4 simulation, it has been possible to measure the calibration constant $(37.2 \pm 0.1^{\text{stat}} \pm 0.3^{\text{syst}}) \text{ mV/MeV}$ where the systematic error takes into account the uncertainty in the telescope acceptance and in the fitting model. Thus, the response of PIPS detector to muons is compatible with the one obtained with γ -rays and also the measured energy resolution, $\sigma \simeq 14 \text{ keV}$ is compatible with previous measurements.

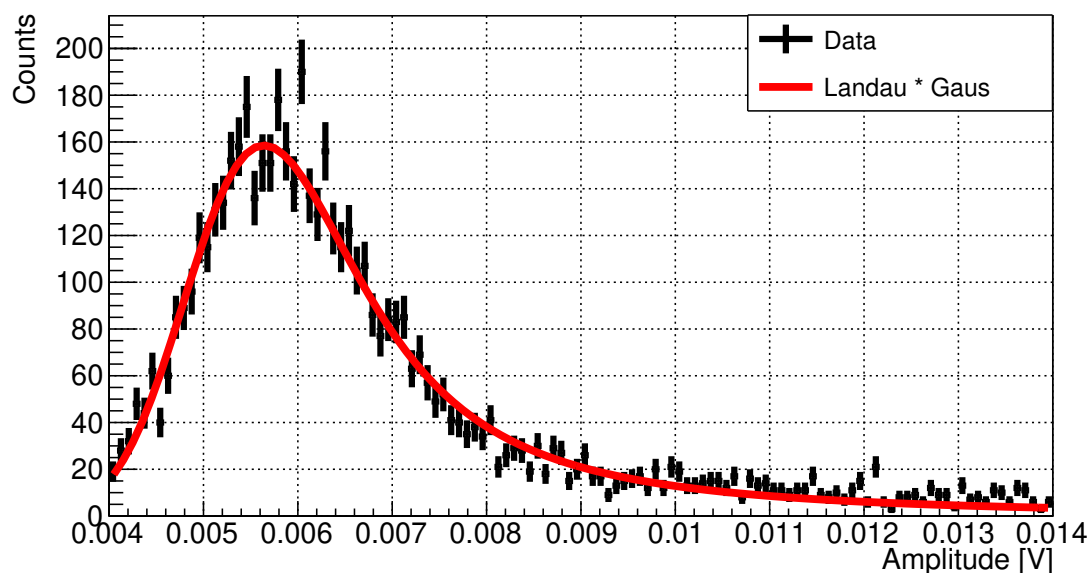


Figure 9. Results from the calibration with muons. In the picture the histogram of the amplitude of the PIPS voltage signals (points). Red line is a fit with a Landau distribution convoluted with a gaussian.

6.3. Characterisation with α -particles emitted by ^{241}Am

Calibration with gamma-rays (producing recoiling electrons) is obtained with particles with a relatively small dE/dx , near the MIP (Minimum Ionising Particle) point. An important test for LEM

silicon detectors is the calibration with slow alpha particles, having a much larger dE/dx . In particular, we use the alpha particles emitted by the calibration source of ^{241}Am . The alpha particles emitted by ^{241}Am have an energy 5.49/5.44/5.4 MeV (mostly). Our ^{241}Am calibration source is manufactured with a very thin layer holding the Am isotope, therefore alpha particles can escape the source losing just a fraction of their energy within the source and in the air; the resulting energy distribution has a spread in angle and energy. However, exploiting that few MeV alpha particles lose roughly $\sim 1 \text{ MeV/cm}$ in air, it is possible to measure the variation of the residual alpha particle energy as a function of the distance between the Ion-Implanted Silicon Detector and the source.

In Figure 10, signals induced by alpha particles are displayed. Signals are acquired with the LeCroy WaveSurfer 3000 oscilloscope. The silicon detector is provided by the manufacturer with a built-in Charge Sensitive Amplifier (CSA) and a shaper. As we can deduce, the width of the signal at 50% of the peak value is about 100 ns. This allows us to say that the maximum particle rate that is possible to acquire with this detector, without a pile-up occurring, is below a few MHz.

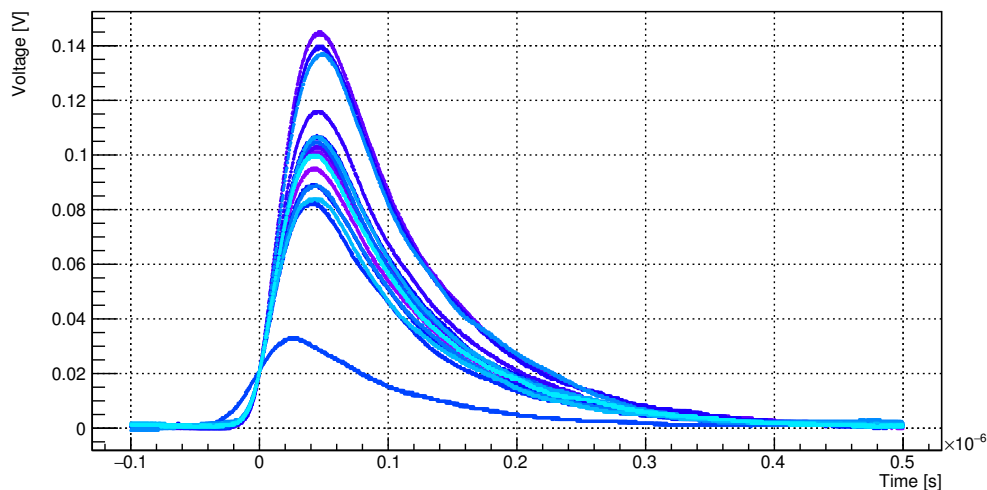


Figure 10. The signal induced by ^{241}Am alpha particles in the PIPS detector. A LeCroy WaveSurfer 3000 oscilloscope has been used as a DAQ.

In Figure 11, it is possible to see different alpha particle amplitude spectra measured for different distances from the ^{241}Am source. To minimise the alpha particle angular spread we apply a collimator made by a tape with a small central hole in front of the source. Unfortunately, the exact thickness (a few mm) of the air gap within the source itself is not known (we did not exactly measure it to avoid damaging the thin source). Moreover, also the source material composition and exact geometry of the Am isotope distribution within the source material is unknown. This adds an unknown, positive, air-equivalent offset to our measured distance.

The measured alpha particle energy distributions are not gaussian and present long tails at low energies, this is due to different energy loss in air and source due to different inclination and multiple scattering but also detector effects can contribute [39]. For this reason, we fitted those spectra with a gaussian curve modified with the addition of a linear tail function. After the fitting procedure we provided an estimation of mean value of the gaussian μ . From Figure 11, it is also possible to estimate the derivative of the signal amplitude as a function of the distance, which is a quantity proportional to the stopping power $\partial E/\partial x$. However, the stopping power is a function of the particle's energy. As a consequence, a quantitative calibration from Figure 11 is challenging because of the unknown air-equivalent distance offset. Thus, by measuring $\frac{1}{\langle E \rangle} \frac{\partial E}{\partial x}$, we can remove the dependence on the calibration constant from voltage to energy. This quantity is still dependent on the true particle kinetic energy. Therefore, by assuming an hypothesis on the alpha particle calibration constant, it is possible

to compare the measurement $\frac{1}{\langle E \rangle} \frac{\partial E}{\partial x}$ with the expectation from a Monte-Carlo simulation of alpha particles crossing the air.

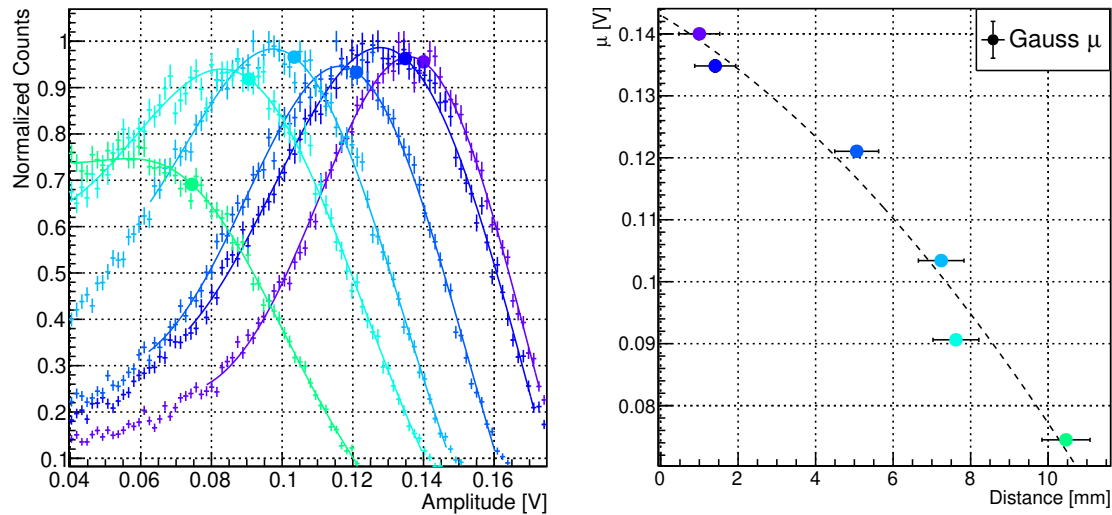


Figure 11. Left: α particle amplitude spectra collected with the PIPS detector (colour encodes the distance). Right: signal amplitude (μ round points) as a function of the distance between the ^{241}Am source and the PIPS, with an unknown offset.

As an example, in Figure 12, the (x-axis) energy estimation adopts the calibration constant measured with gamma sources (described in 6.1).

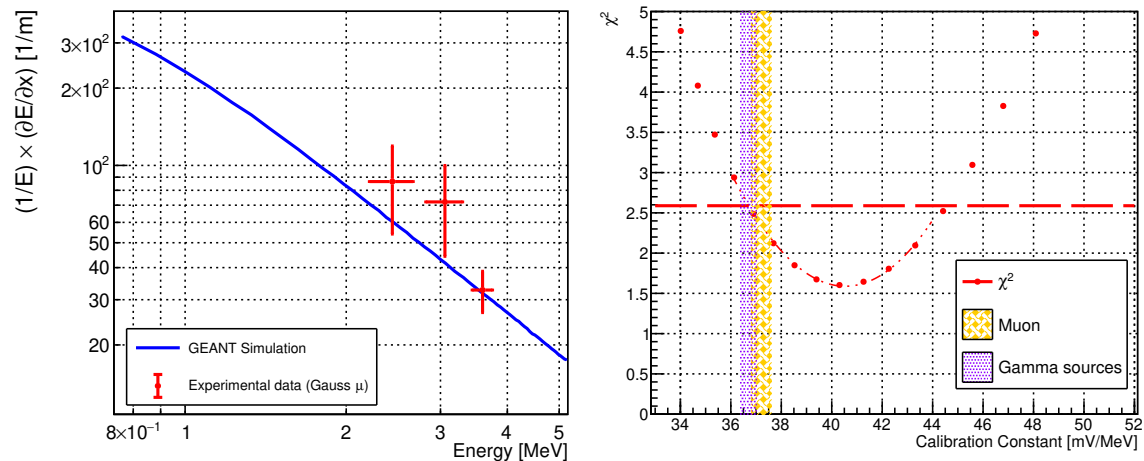


Figure 12. Left: PIPS (fully depleted silicon detector) characterisation with alpha particles. On the x-axis the mean energy detected by the PIPS detector assuming the detector calibration constant as measured with gamma-rays. Right: χ^2 computed for the different hypothesis of the detector calibration constant for alpha particles. Pink shaded area depicts the calibration constant measured with gamma-rays and yellow shaded area depicts the one measured with atmospheric muons.

Finally, scanning different values for the calibration constant, we can measure this parameter for alpha particles by the computation of the χ^2 for each hypothesis as shown in right panel of Figure 12. We obtain that $\chi^2_{min} \sim 1.6$ is reached at 40.5 mV/MeV. From the standard confidence interval evaluation we deduce that the obtained result is compatible with the calibration constant obtained for gamma-ray and muons. Thus, the detector behaves as expected/desired and no evidence

for a detector response depending on particle dE/dx is found. Summarising the result of this analysis, the measured calibration constant for alpha particles is (42 ± 4) mV/MeV.

7. Conclusions and Outlooks

In this work, we described the Low-Energy Module (LEM): a compact particle spectrometer, suitable for a CubeSat, for measurement of the differential flux of low-energy particles in the lower magnetosphere. Here, it is worth summarising the structure of the LEM. To avoid a bulky and heavy detector, we designed an active collimator based on a thin aluminium shield followed by an anti-coincidence detector. The drilled aluminium shield protects the drilled ACD, made of a plastic scintillator, from the large flux of very low energetic electrons in LEO. The holes in the aluminium and in the ACD are used to select a known direction of the particles with an angular resolution of 6° – 7° . The LEM field of view is $60^\circ \times 60^\circ$, monitoring 16 directions in the sky at the same time. The particle identification relies on a series of 16 $\Delta E - E$ modules, based on the PIPS and CdZnTe detectors, placed below each collimator channel. An additional layer of plastic scintillator at the bottom is added as a veto to identify non-contained particles. This detection concept seems to be promising for Particle Identification (PID) at low energies and for gamma/X-ray detection.

In the last part of this work we provided the characterisation procedure we adopted for testing one candidate sensor: the PIPS-APCAM25 $500\mu\text{m}$ fully depleted silicon detector manufactured by Mirion. The sensors came with an embedded preamplifier also object of this study. The calibration and characterisation procedure is essential for tuning the MonteCarlo simulation and deducing the performances of the instrument payload.

We calibrated the sensor by exploiting γ -ray sources and determine the energy resolution of $(11.1 \pm 0.4^{\text{stat}} \pm 1.2^{\text{syst}})$ keV. The signal amplitude calibration constant determined was (36.7 ± 0.3) mV/MeV, no dependence from the particle nature was observed. In particular, the result determined by this traditional calibration is compatible with the result obtained with muon detection $(37.2 \pm 0.1^{\text{stat}} \pm 0.3^{\text{syst}})$ mV/MeV and alpha particles (42 ± 4) mV/MeV. The characterisation also confirmed the absence of a signal's dependence on the reverse-bias voltage applied (above the depletion voltage).

Additional tests on prototypes of CZT detectors are planned in the future at the INFN-TIFPA laboratory.

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Abbreviations

The following abbreviations are used in this manuscript:

ACD	Anti-Coincidence Detector
CSA	Charge Sensitive Amplifier
CSES	China Seismo-Electromagnetic Satellite
CZT	Cadmium Zinc Telluride (CdZnTe)
GEANT	Geometrtty And Tracking
GRB	Gamma-Ray Burst
LAIM	Lithosphere Atmosphere Ionosphere Magnetosphere
EM	Low-Energy Module
LEO	Low Earth Orbit
LYSO	Lutetium-Yttrium OxyorthoSilicate
MILC	Magnetosphere Ionosphere Lithosphere Coupling
MIP	Minimum Ionising Particle
MPV	Most Probable Value
PID	Particle Identification
PIPS	Passivated Implanted Planar Silicon

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