

Article

GRAVITATION IN UNIFIED SCALAR FIELD THEORY

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Abstract: The scalar field of space-time film is considered as unified fundamental field. The field model under consideration is the space-time generalization of the model for a two-dimensional thin film. The force and metrical interactions between solitons are considered. These interactions correspond to the electromagnetic and gravitational interactions respectively. The metrical interaction and its correspondence to the gravitational one are considered in detail. The practical applications of this approach are discussed.

Keywords: Unified field; space-time film; soliton-particle; gravitation; dark matter; gravitational screening

1. Introduction

1.1. Unified fundamental field

The concept of unified fundamental field exists for a long time sufficiently. According to this concept, matter in all its multiformity can be considered as a solution of some nonlinear field model.

In particular, all elementary particles must be represented by soliton solutions of the model. In this case their interactions are the consequence of the field model nonlinearity.

The nonlinearity violates the superposition property for solutions of the appropriate linearized model. According to this property the sum of solutions is also solution. The violation of the superposition property is interpreted as the interaction of particles.

It is reasonable that the interaction of big material objects containing many particles is explained by the same way.

Attempts to propose models of the unified field theory have been made by many well-known researchers. It is known that the creator of general relativity theory A. Einstein was adherent of this concept. He tried to create an unified field theory for the rest of his life.

1.2. Attempts to create an unified field theory

Electromagnetic field satisfied nonlinear equations was considered as unified field. First of all we must notice two known nonlinear electrodynamics model, namely, Mie [14] and Born – Infeld models [2].

Einstein considered, in particular, the field of metrical tensor for curved space-time as unified field [11]. In this connection nonsymmetrical metric was also considered [12].

Heisenberg tried to build an unified theory of spinor field [13].

All these attempts were not completely successful. In particular, within the framework of these theories, no solution has been obtained that corresponds with a sufficient degree of realism to any elementary particle. It should be noted that this problem is extremely difficult mathematically.

1.3. The question as to the tensor rank of the field

It is evident that the question as to the tensor rank of an unified field is highly important.

In the above-mentioned attempts to construct an unified field theory, the choice of the tensor rank of the field was based on certain correspondence considerations.

For example, the nonlinear electrodynamics models generalized the linear electrodynamics which was successful in certain limits.

The variants of Einstein's unified theories generalized his own gravitational theory which was successful in certain limits.

Heisenberg's theory of an unified spinor field was based on Dirac's linear theory which was successful in certain limits.

1.4. The force and metrical interactions of particles-solitons

As a natural criterion for the acceptability of the unified field model, we should consider the possibility in its framework for description of the two long-range interactions of material objects, namely electromagnetism and gravitation.

A number of the author's works were devoted to the unification of electromagnetism and gravitation in the framework of Born – Infeld nonlinear electrodynamics (see, for example, [3–6])

In this case, the electromagnetic interaction of particles-solitons is a consequence of the integral conservation law for energy-momentum. This type of interaction is called the force one.

The description of gravity in the framework of nonlinear electrodynamics is based on the effect of induced space-time curvature by a weak field of distant particles-solitons at the location of the test particle. This type of interaction is called the metrical one.

1.5. Scalar field

Force and metric interactions of particles-solitons are inherent in all nonlinear field models that are invariant under shifts and rotations in four-dimensional space-time.

Thus we can consider a simpler model of a scalar unified field, since it also allows us to describe electromagnetism and gravity [7].

In particular, the tensor character of the electromagnetic field is due to the determination of the integral force through the integral over a closed surface surrounding the test particle under the force interaction [8].

It should be emphasized that the electromagnetic field of antisymmetric tensor of the second rank appears at a certain point when a test particle is placed there. In the absence of such a particle, only a configuration of the scalar field is present near this point, but it generates the electromagnetic interaction with any charged particle.

2. The unified field

2.1. Space-time film

We consider the following generally covariant world volume action and the appropriate variational principle [9]:

$$\mathcal{A} = \int_{\bar{V}} \sqrt{|\mathfrak{M}|} (dx)^4 = \int_{\bar{V}} \mathcal{L} d\bar{V}, \quad \delta\mathcal{A} = 0, \quad (1a)$$

where $\mathfrak{M} \doteq \det(\mathfrak{M}_{\mu\nu})$, $(dx)^4 \doteq dx^0 dx^1 dx^2 dx^3$, \bar{V} is space-time volume, $d\bar{V} \doteq \sqrt{|\mathfrak{m}|} (dx)^4$ is four-dimensional volume element, $\mathfrak{m} \doteq \det(\mathfrak{m}_{\mu\nu})$,

$$\mathfrak{M}_{\mu\nu} = \mathfrak{m}_{\mu\nu} + \chi^2 \frac{\partial\Phi}{\partial x^\mu} \frac{\partial\Phi}{\partial x^\nu}, \quad \mathcal{L} \doteq \sqrt{\left| 1 + \chi^2 \mathfrak{m}^{\mu\nu} \frac{\partial\Phi}{\partial x^\mu} \frac{\partial\Phi}{\partial x^\nu} \right|} \quad (1b)$$

$\mathfrak{m}_{\mu\nu}$ are components of metric tensor for flat four-dimensional space-time, Φ is scalar real field function, χ is dimensional constant. The Greek indices take values $\{0, 1, 2, 3\}$. The tensor $\mathfrak{M}_{\mu\nu}$ can be called the world tensor.

The model (1) can be considered as a relativistic generalization of the appropriate expression for the mathematical model of two-dimensional minimal thin film in the tree-dimensional space of our everyday experience.

2.2. Energy-momentum density tensor

Customary method gives the following canonical energy-momentum density tensor of the model in Cartesian coordinates

$$\bar{\mathfrak{F}}^{\mu\nu} = \frac{1}{4\pi} \left(\frac{\Phi^\mu \Phi^\nu}{\mathcal{L}} - \frac{\mathfrak{m}^{\mu\nu}}{\chi^2} \mathcal{L} \right), \quad \Phi^\alpha \doteq \underline{\mathfrak{m}}^{\alpha\beta} \frac{\partial\Phi}{\partial x^\beta}. \quad (2)$$

where $\underline{\mathfrak{m}}^{\mu\nu}$ is the constant diagonal metrical tensor for flat space-time with signature $\{+, -, -, -\}$ or $\{-, +, +, +\}$. As we see, the canonical tensor is symmetrical.

To obtain finite integral characteristics of solutions in infinite space-time we introduce the regularized energy-momentum density tensor with the following formula:

$$\mathfrak{F}^{\mu\nu} = \bar{\mathfrak{F}}^{\mu\nu} - \overset{\infty}{\mathfrak{F}}^{\mu\nu}. \quad (3)$$

where $\overset{\infty}{\mathfrak{F}}^{\mu\nu}$ is a regularizing symmetrical energy-momentum density tensor. Here we will use the constant regularizing tensor

$$\overset{\infty}{\mathfrak{F}}^{\mu\nu} = -\frac{1}{4\pi\chi^2} \underline{\mathfrak{m}}^{\mu\nu}. \quad (4)$$

In general case we can take a symmetrical tensor satisfying the differential conservation law as the regularizing tensor. A special choice of this tensor can provide the convergence of energy integral for a certain class of solutions.

2.3. Equation of space-time film

The variational principle (1) gives the following model field equation in Cartesian coordinates [9]:

$$\left(\underline{\mathfrak{m}}^{\mu\nu} \left(1 + \chi^2 \underline{\mathfrak{m}}_{\sigma\rho} \Phi^\sigma \Phi^\rho \right) - \chi^2 \Phi^\mu \Phi^\nu \right) \frac{\partial^2 \Phi}{\partial x^\mu \partial x^\nu} = 0, \quad (5)$$

This field equation can be written in the following remarkable form [7]:

$$\tilde{\mathfrak{m}}^{\mu\nu} \frac{\partial^2 \Phi}{\partial x^\mu \partial x^\nu} = 0, \quad (6a)$$

$$\tilde{\mathfrak{m}}^{\mu\nu} \doteq -4\pi\chi^2 \bar{\mathfrak{F}}^{\mu\nu}. \quad (6b)$$

where $\bar{\mathfrak{F}}^{\mu\nu}$ is the canonical energy-momentum density tensor (2). Here we introduce the effective metric $\tilde{\mathfrak{m}}^{\mu\nu}$ which will be considered below.

Equation (6) transforms to ordinary linear wave equation with $\chi = 0$:

$$\underline{m}^{\mu\nu} \frac{\partial^2 \Phi}{\partial x^\mu \partial x^\nu} = 0. \quad (7)$$

2.4. Effective metric and curved space-time

The characteristic equation of the model has the following form which is obtained directly from (6a):

$$\tilde{m}^{\mu\nu} \tilde{k}_\mu \tilde{k}_\nu = 0, \quad \tilde{k}_\mu \doteq \frac{\partial \mathcal{S}}{\partial x^\mu}, \quad (8)$$

Equation $\mathcal{S} = 0$ gives a three-dimensional characteristic hypersurface of the field model in four-dimensional space-time.

As is known, the problem of propagation of a weak high-frequency quasi-plane wave with a given background field in a nonlinear model gives the dispersion relation for the components of the wave four-vector that coincides with the characteristic equation [6]. The dispersion relation or a dependence between the frequency of the wave and its three-dimensional wave vector defines the wave front propagation.

Thus the components \tilde{k}_μ in equation (8) can also be considered as the components of the wave four-vector of a high-frequency quasi-plane wave propagating in an effective curved space-time with the metric $\tilde{m}^{\mu\nu}$.

Equation (8) sets the motion of such wave packets that, in the absence of a background field, would move at the speed of light. Thus this equation determines the trajectories of massless particles. For the trajectory of a massive particle, the considered approach gives the equation of a geodesic line in an effective curved space-time with the metric $\tilde{m}^{\mu\nu}$. This theme is considered below in section 4.

3. Solitons-particles

3.1. Solitons are material particles

Solitons are spatially localized solutions of nonlinear field models. Solitons-particles are soliton solutions of the unified field model corresponding to elementary particles.

Another name for soliton is solitary wave. But linear field models can also have the solutions in the form of solitary waves. Because of linearity of the models these solutions can have an arbitrarily small amplitude. Thus we shall call these solitary waves the weak solitons.

Not long ago, a class of exact soliton solutions propagating with the speed of light was obtained for the model of space-time film [9]. A subclass of these solutions has properties which allow to correlate it with photons. Other solutions of the found class of exact solutions may relate to different neutrinos.

At the present time, among exact soliton solutions for massive particles we know the simplest spherically symmetrical soliton in intrinsic coordinate system or spheron [8]. It has a rest energy and electrical charge but does not have angular momentum or spin.

There is a defined progress in the finding of approximate soliton solutions of the toroidal configuration containing an oscillating part [10]. These solutions can also have a spin in addition to the rest mass and charge.

3.2. Oscillating parts of solitons-particles

First of all the existence of oscillating parts in solitons-particles is necessary for the description of observable wave properties for elementary particles.

But in the framework of the unified field theory under consideration the oscillating parts are necessary for the description of real gravitation. We shall discuss this question below.

When we move away from the localization region of a soliton-particle, the field of space-time film satisfies approximately the linear wave equation (7).

Elementary oscillating solutions of the linear wave equation in the spherical coordinate system decrease in amplitude as r^{-1} . Let us take the following simple solution as example:

$$\overset{\circ}{\Phi} = \frac{\overset{\circ}{a}}{r} \sin(\omega r) \sin(\omega \overset{\circ}{x}^0), \quad (9a)$$

where the point under symbols denotes the belonging to the intrinsic coordinate system.

According to the definition given in the previous subsection, this solution is a weak soliton.

The asymptotic form such as we have in (9a) leads to the divergence of energy integral at infinity.

However, the investigation of toroidal configurations [10] shows that there can be solitons with oscillating part without the asymptotic form of type (9a) whose energy is finite. But we can suppose in this case that the wave mode of type (9a) will appear for interacting solitons.

Thus we assume that a soliton-particle in the intrinsic coordinate system has both static and oscillating parts. The oscillating part is a standing wave having perhaps sufficiently complicated configuration. Using a space-time rotation we can obtain a moving soliton from the rest one. In this case the standing wave transforms to moving one. The moving soliton is obtained with the help of the following substitution for the intrinsic coordinates of the soliton:

$$\overset{\circ}{x}^\mu = L_{\nu}^{\mu} \overset{\circ}{x}^{\nu}, \quad (9b)$$

where $L_{\nu}^{\mu} L^{\rho\nu} = L_{\nu}^{\mu} L^{\nu\rho} = \underline{m}^{\mu\rho}$, L_{ν}^{μ} is a matrix of space-time rotation, in particular, Lorentz transformations, $\{\overset{\circ}{x}^{\nu}\}$ is the coordinate system in which the soliton is moving.

We have the following dispersion relation for the wave four-vector components of this moving wave:

$$\left| \underline{m}^{\mu\nu} k_{\mu} k_{\nu} \right| = \omega^2, \quad \omega \doteq -k_0, \quad (10)$$

where ω is the angular frequency of the standing wave in intrinsic coordinate system of the soliton, $\{\omega, k_i\}$ are the frequency and wave vector components of the traveling wave.

As can be verified, the group velocity of the traveling wave corresponding to the dispersion relation (10) coincides with the velocity parameter $\overset{\circ}{v}$ of the space-time rotation (9b). The frequency ω is called a rest frequency of the soliton.

4. Gravitation

4.1. Gravitation as the metrical interaction of solitons

The effect of induced gravitation in nonlinear electrodynamics is the subject of a sufficient number of works by the author, including an article in the encyclopedia [4] and a monograph [6].

Now this analysis is applied to the space-time film model with insignificant modification.

Let us consider the problem of propagation of a small amplitude wave with a rest frequency against a background given field of distant solitons-particles.

For distinctness we take the simple solution of the linear wave equation (9a). This solution has the form of a standing wave in the intrinsic coordinate system. It has a finite amplitude at the coordinate origin $\overset{\circ}{a}$ which can be taken sufficiently small.

Let us consider the sum of the background field $\tilde{\Phi}$ generated by distant solitons and a moving fast-oscillating weak soliton $\overset{\circ}{\Phi}$ of type (9) but with a slow time-dependent velocity parameter $\overset{\circ}{v}$ and a small constant amplitude $\overset{\circ}{a}$:

$$\Phi = \tilde{\Phi} + \overset{\circ}{\Phi}, \quad (11a)$$

$$|\omega| \gg \max \left| \frac{\partial \tilde{\Phi}}{\partial x^\mu} \right| / \max |\tilde{\Phi}|, \quad \left| \frac{d\overset{\circ}{v}}{dx^0} \right| \ll |\omega|, \quad \overset{\circ}{a} \ll \max |\tilde{\Phi}|. \quad (11b)$$

Here we consider the weak soliton with a constant amplitude $\overset{\circ}{a}$ as some approximation for a soliton-particle whose amplitude is defined by an exact solution. If there is an additional field of remote solitons $\tilde{\Phi}$, the soliton solution is modified, and its amplitude could also be changed. However, it is natural to assume that the soliton-particle has a maximum amplitude which is significantly greater than the field of distant solitons-particles. Then the weak field of distant particles will not significantly affect the amplitude of the considered soliton-particle. Thus if we consider the weak soliton instead of the soliton-particle, we actually investigate the influence of the distant solitons field $\tilde{\Phi}$ to the part of soliton-particle which is sufficiently far from its center and has a small amplitude. The movement of this weak part should direct the entire soliton-particle, since it is a modified exact solution.

We substitute this sum (11) in the equation of space-time film (6). Here we suppose an averaging of the effective metric $\tilde{m}^{\mu\nu}$ (6b) with the background field $\tilde{\Phi}$ over a defined space-time localization region for the weak soliton $\overset{\circ}{\Phi}$. Then we obtain the following equation:

$$\tilde{m}^{\mu\nu} \frac{\partial^2 \overset{\circ}{\Phi}}{\partial x^\mu \partial x^\nu} = 0. \quad (12)$$

where the averaging effective metric $\tilde{m}^{\mu\nu}$ depends on the derivatives of the background field: $\tilde{m}^{\mu\nu} = \tilde{m}^{\mu\nu}(\tilde{\Phi})$. Here below we use the designation $\tilde{m}^{\mu\nu}$ for the averaging effective metric $\tilde{m}^{\mu\nu}$ in an effort to simplify the designations.

We consider that the averaging background effective metric $\tilde{m}^{\mu\nu}$ is almost constant in the space-time localization region of the weak soliton $\overset{\circ}{\Phi}$. In this area, let us find a coordinate transformation which reduces the equation (12) to ordinary wave one (7).

Let us have the following relations for the coordinate differentials:

$$dx^\mu = \hat{X}^\mu_\nu dx^\nu, \quad dx^\mu = \check{X}^\mu_\nu dx^\nu, \quad (13)$$

where the matrix \hat{X}^μ_ν and \check{X}^μ_ν are mutually inverse and satisfy the relations

$$\tilde{m}^{\mu\nu} \hat{X}^\sigma_\mu \hat{X}^\rho_\nu = \underline{m}^{\sigma\rho}, \quad \tilde{m}^{\mu\rho} = \check{X}^\mu_\zeta \check{X}^\rho_\xi \underline{m}^{\zeta\xi}. \quad (14)$$

Then the solutions of equation (12) in the limited area under consideration has the form

$$\overset{\circ}{\Phi} = \overset{\circ}{\Phi}(x^\sigma(x)). \quad (15)$$

where $\overset{\circ}{\Phi}(x^\sigma)$ is a solution of the linear wave equation (7), in particular, the weak soliton (9a). Here the functions $x^\sigma(x^\delta)$ are defined by the transformation (13).

The averaged background effective metric $\tilde{m}^{\mu\nu}$ calculated in an expanded four-dimensional space defines a Riemann space that is not generally flat. Also, in general, it is not possible to find the coordinate transformation $x^\sigma = x^\sigma(x^\delta)$ satisfying (13) and (14) everywhere.

But we can write this transformation approximately in the following form:

$$x^i = \hat{X}_j^i(x^0) \left(x^j - \mathfrak{x}^j(x^0) \right), \quad (16a)$$

$$x^0 = \hat{x}^0(x^\nu) = \int \bar{k}_\mu dx^\mu, \quad \bar{k}_\mu(x^\nu) \doteq \hat{X}_{,\mu}^0(x^\nu). \quad (16b)$$

where $\mathfrak{x}^j(x^0)$ is a position of energy center for the weak soliton at the time x^0 , \bar{k} is called a normalized wave vector such that

$$\frac{\partial \bar{k}_\mu}{\partial x^\nu} = \frac{\partial \bar{k}_\nu}{\partial x^\mu}. \quad (16c)$$

Taking into account the transformation (16) we obtain that ωx^0 is a phase of the wave $\hat{\Phi}$. Then we have the following its wave vector components:

$$k_\mu = \omega \bar{k}_\mu. \quad (17)$$

It should be emphasized that the substitution of type (16b) for x^0 to the solution (9a) is the known approach to solving the problem of wave propagation in an inhomogeneous medium [15]. In this approach, the wave phase is considered as an unknown coordinate function, and the substitution of the modified wave solution into the equation gives a dispersion relation for this wave.

Using (16b) and (14), we have the following dispersion relation for the weak soliton:

$$\tilde{\mathfrak{m}}^{\mu\nu} \bar{k}_\mu \bar{k}_\nu = \underline{\mathfrak{m}}^{00}, \quad (18)$$

where $\tilde{\mathfrak{m}}^{\mu\nu} = \tilde{\mathfrak{m}}^{\mu\nu}(x^\rho)$, $\bar{k}_\mu = \bar{k}_\mu(x^\rho)$.

Now let us obtain a trajectory of energy center for the weak soliton $\mathfrak{x}^j = \mathfrak{x}^j(x^0)$. We use the intrinsic time $x^0(x^\mu)$ (16b) at the point \mathfrak{x}^μ as a parameter of movement \mathfrak{s} . Taking also $\mathfrak{x}^0 = x^0$ we have

$$\mathfrak{s} \doteq x^0(\mathfrak{x}^\mu) = \mathfrak{s}(x^0), \quad (19a)$$

$$\mathfrak{u}^\mu \doteq \frac{d\mathfrak{x}^\mu}{d\mathfrak{s}} = \hat{X}_{,0}^\mu(\mathfrak{x}^\nu). \quad (19b)$$

where \mathfrak{u}^μ is the four velocity. We have the last relation in (19b) in accordance to (13) and (19a).

The definitions \mathfrak{u}^μ (19b) and \bar{k}_ν (16b) with (14) gives also the following relations at the point $\{\mathfrak{x}^\mu\}$:

$$\bar{k}_\mu \mathfrak{u}^\mu = 1, \quad (20a)$$

$$\mathfrak{u}^\mu = \underline{\mathfrak{m}}^{00} \tilde{\mathfrak{m}}^{\mu\nu} \bar{k}_\nu. \quad (20b)$$

Let us introduce the inverse tensor $\tilde{\mathfrak{m}}_{\mu\nu}$ to the tensor $\tilde{\mathfrak{m}}^{\mu\nu}$:

$$\tilde{\mathfrak{m}}_{\mu\nu} \tilde{\mathfrak{m}}^{\nu\rho} = \delta_\mu^\rho. \quad (21)$$

We have from (20) and (21) the following relation at the point $\{\mathfrak{x}^\mu\}$:

$$\tilde{\mathfrak{m}}_{\mu\nu} \mathfrak{u}^\mu \mathfrak{u}^\nu = \underline{\mathfrak{m}}^{00}. \quad (22)$$

This leads to the well-known expression in general relativity theory:

$$d\mathfrak{s}^2 = \frac{1}{\underline{\mathfrak{m}}^{00}} \tilde{\mathfrak{m}}_{\mu\nu} d\mathfrak{x}^\mu d\mathfrak{x}^\nu = \left| \tilde{\mathfrak{m}}_{\mu\nu} d\mathfrak{x}^\mu d\mathfrak{x}^\nu \right|. \quad (23)$$

We have also from (20) the following relation at the point x^μ

$$\bar{k}_\mu = \frac{1}{\underline{m}^{00}} \check{m}_{\mu\nu} u^\nu \quad (24)$$

One must note that although u^μ and \bar{k}_μ are defined as components of second-order matrix, they are actually four-vectors.

Indeed, if we consider the problem in another coordinate system $\{x'^\mu\}$ then we have

$$ds \doteq dx^0 = \hat{X}_{,\mu}^0 dx^\mu = \hat{X}'_{,\mu}{}^0 dx'^\mu . \quad (25)$$

This means that the intrinsic time x^0 is not changed for such transformations or it behaves as scalar. Thus \bar{k}_μ is a four-vector according to the definition (16b).

According to (16a), the three-dimensional coordinate system $\{x^i\}$ is moving coupled with the localization region of the weak soliton. Thus we have $\dot{x}^i = 0$ and $dx^i = 0$ at the point $\{x^\mu\}$ and

$$dx^\mu = \check{X}_{,0}^\mu(x^\nu) dx^0, \quad dx'^\mu = \check{X}'_{,0}{}^\mu(x'^\nu) dx^0 . \quad (26)$$

Therefore u^μ are a four-vector according to definition (19b).

Now let us obtain the trajectory equation for $x^\mu(s)$. Differentiation of dispersion relation (18) with respect to certain coordinate x^ρ with consideration of relations (16c) and (20) gives the following equation at the point x^μ :

$$\begin{aligned} \frac{\partial \check{m}^{\mu\nu}}{\partial x^\rho} \bar{k}_\mu \bar{k}_\nu + 2 \underline{m}^{00} u^\nu \frac{\partial \bar{k}_\rho}{\partial x^\nu} &= 0 \\ \implies \frac{\partial \check{m}^{\mu\nu}}{\partial x^\rho} \bar{k}_\mu \bar{k}_\nu + 2 \underline{m}^{00} \frac{d\bar{k}_\rho}{ds} &= 0 . \end{aligned} \quad (27)$$

Substituting (24) into (27) and using (21) we obtain the trajectory equation in the following form:

$$\frac{d\mathbf{u}^\mu}{ds} + \tilde{\Gamma}_{\nu\rho}^\mu u^\nu u^\rho = 0 , \quad (28a)$$

where

$$\tilde{\Gamma}_{\nu\rho}^\mu \doteq \frac{1}{2} \check{m}^{\mu\delta} \left(\frac{\partial \check{m}_{\delta\nu}}{\partial x^\rho} + \frac{\partial \check{m}_{\delta\rho}}{\partial x^\nu} - \frac{\partial \check{m}_{\nu\rho}}{\partial x^\delta} \right) . \quad (28b)$$

As can be seen, this equation is the geodesic line one for the introduced effective Riemann space with the metric $\check{m}^{\mu\nu}$.

Thus a weak soliton with rest frequency and small constant amplitude under the influence of distant solitons behaves as massive particle in gravitational field.

We can assume that the obtained equation (28) describes also the movement of a soliton-particle under the influence of distant solitons-particles.

4.2. Newtonian potential

The geodesic line equation (28) in zero speed approximation contains the component \check{m}_{00} only. Let us write it in the following form:

$$\check{m}_{00} = \pm (1 - 2\varphi) , \quad \varphi \doteq -2\pi \chi^2 \mathcal{E} \quad (29)$$

where φ is the scalar potential of the gravitational field, $\bar{\mathcal{E}}$ is an averaging energy density for the field of distant solitons. Here the averaging is performed over a space-time volume including a localization region of the soliton and a relevant time interval

In order to have the real gravitation in our consideration we must obtain the following expression for the gravitational potential:

$$\varphi = -\gamma \frac{\tilde{m}}{R}, \quad (30)$$

where γ is the gravitational constant, \tilde{m} is a mass for agglomeration of distant solitons-particles, R is a distance from the energy center of this agglomeration.

Thus we must obtain the following asymptotic form for the averaging energy density as $R \rightarrow \infty$:

$$\bar{\mathcal{E}} \sim \frac{1}{R}. \quad (31)$$

4.3. The role of wave background

Let us elucidate how we can obtain the necessary asymptotic form for the averaging energy density (31).

If we consider only static parts of charged solitons-particles ($\Phi \sim r^{-1}$) then the appropriate energy density will decrease apparently as r^{-4} .

Taking into consideration the oscillating parts of interacting solitons-particles with decreasing amplitude as r^{-1} (as, for example, for the weak soliton (9a)) gives the decrease of the averaging energy density as r^{-2} .

To obtain the asymptotic r^{-1} for the potential φ we must also take into account a wave background with almost constant amplitude in space. This wave background must undoubtedly exist in the space where there is a bulk of oscillating solitons-particles.

Then we must represent the field $\tilde{\Phi}$ generating the effective metric in the form of the following sum:

$$\tilde{\Phi} = \tilde{\Phi}^\circ + \tilde{\Phi}^\sim, \quad (32)$$

where $\tilde{\Phi}^\circ$ is the field of distant solitons containing the fast oscillating part with amplitude decreasing as r^{-1} (just as in the weak soliton (9a)), $\tilde{\Phi}^\sim$ is the wave background field with almost constant amplitude.

The substitution of the field (32) to the energy density gives the terms which are proportional to r^{-1} . The following averaging of energy density can give the appropriate term in expression for the potential φ which is proportional to r^{-1} . This expectation seems reasonable because the interaction between the fields $\tilde{\Phi}^\circ$ and $\tilde{\Phi}^\sim$ caused by nonlinearity of the model can give a defined phase synchronization for these fast oscillating fields.

Thus the gravitational constant is defined by two factors

$$\gamma \sim \overset{\circ}{\tilde{\gamma}} \tilde{\gamma}. \quad (33)$$

where $\overset{\circ}{\tilde{\gamma}}$ is a constant of proportionality between an amplitude of the fast oscillating field of distant agglomeration of solitons-particles and its total mass \tilde{m} , $\tilde{\gamma}$ defines the amplitude of the wave background in the considered space-time area.

5. Practical applications

5.1. Possible explanation for the dark matter effect

We assume that the amplitude of the wave background (the factor $\tilde{\gamma}$ in (33)) may vary slightly in space. In this case we must assume that the gravitational constant is not really constant but it can also change slightly. Perhaps the observable effects of so-called dark matter [1] can be explained by a weak spatial dependence of the gravitational constant $\gamma(\mathbf{x})$.

It should also be noted that the concept of gravitational constant is connected with the concept of point mass. But the effective metric in the induced gravitation theory is defined by the wave background and the field of distant solitons-particles. The latter can have various spatial forms that satisfy the model equation asymptotically. The appropriate ideas for explanation of the dark matter effect require further study and separate publication.

5.2. About the possibility of gravitational screening

We have seen that in this approach, the existence of the long-range gravitational potential $\sim r^{-1}$ is associated with the presence of the wave background.

Then if the wave background is cut off or weakened with the help of some method, then the gravitational interaction will also be weakened. In this case we can talk about gravitational screening.

Estimations of possible frequency of the wave background show [6] that the gravitational screening requires a metamaterial from which a gamma-ray mirror can be made.

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Sample Availability: Samples of the compounds are available from the authors.